# FORMATION OF SATELLITES from a tidal disk



#### **Aurélien CRIDA**

& Sébastien CHARNOZ

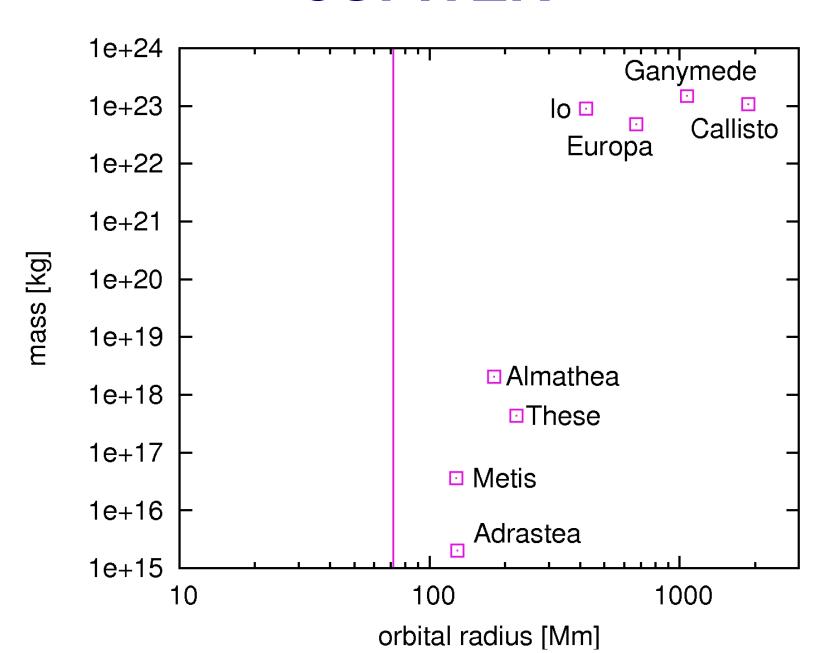




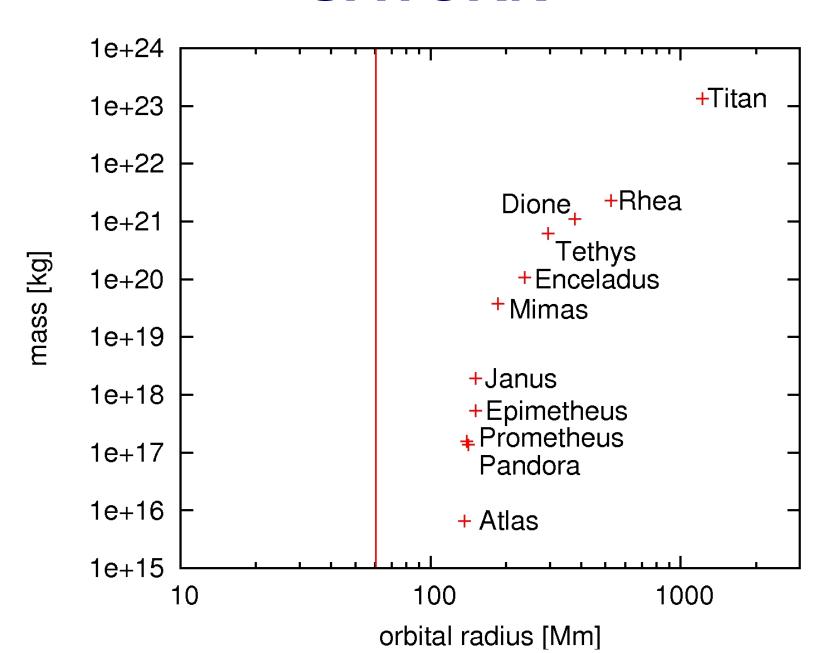




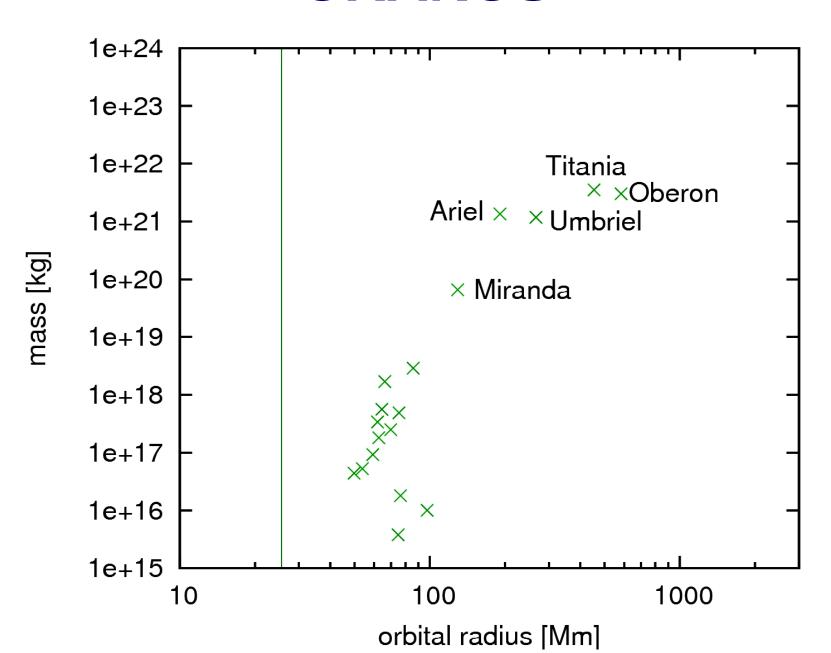
#### **JUPITER**



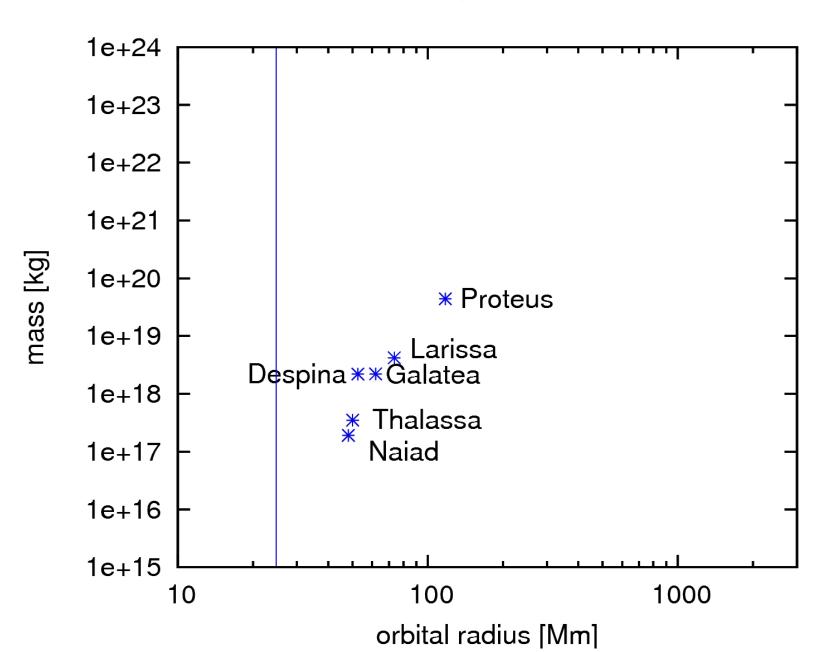
#### **SATURN**



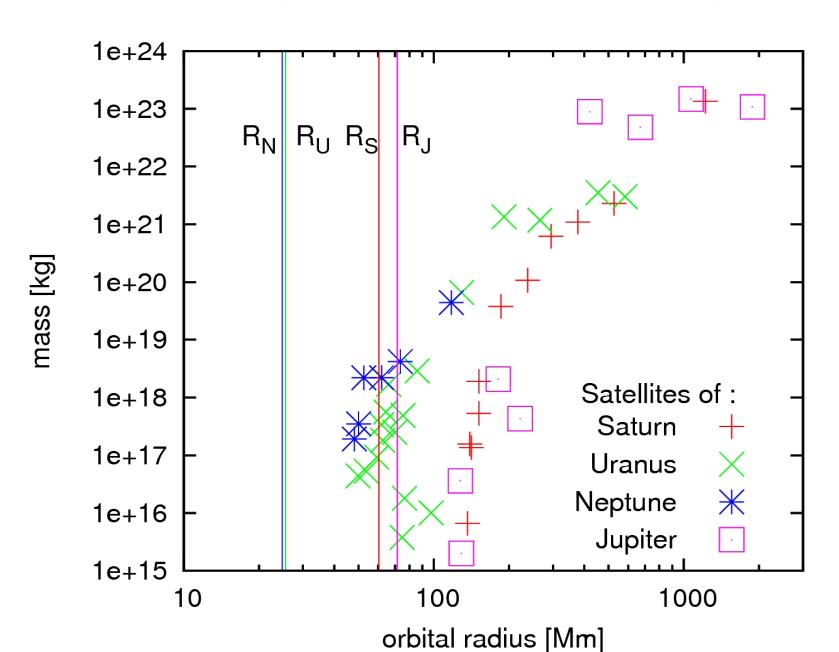
#### **URANUS**



#### **NEPTUNE**



#### **ALL GIANT PLANETS**

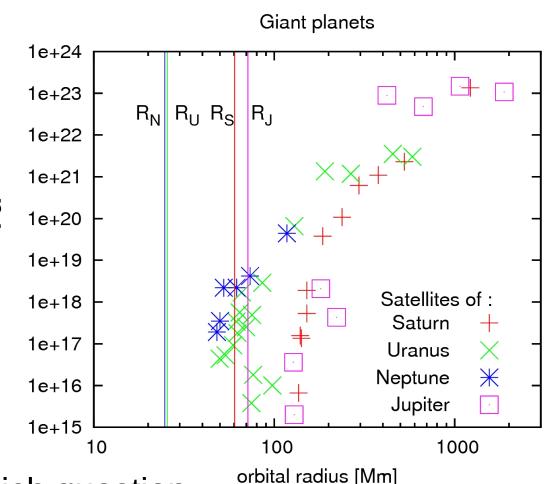


#### INTRODUCTION

Distributions of giant planets' regular satellites:

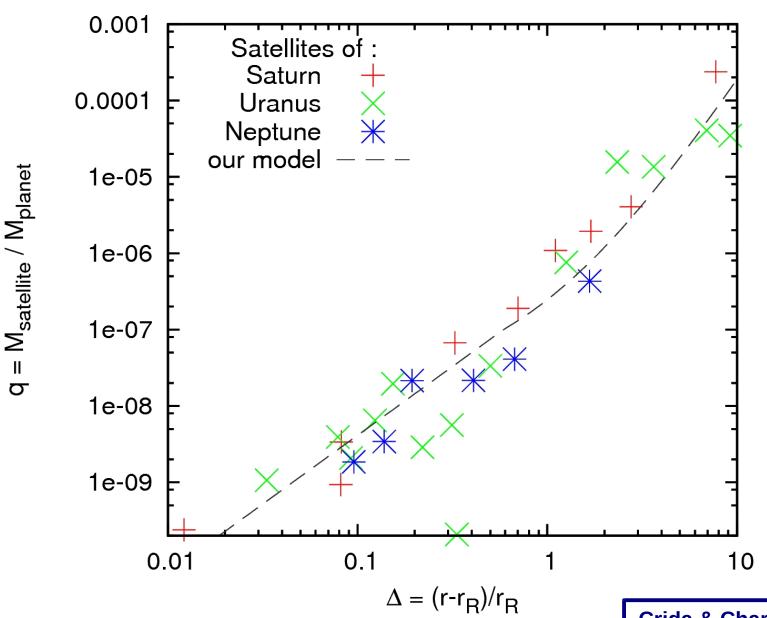
- don't reach the planet
- ranked by mass
- pile-up at a few planetary radii (small bodies)

Why?



It's not a power law, which question the Circum-Planetary Disc model...

#### CONCLUSION

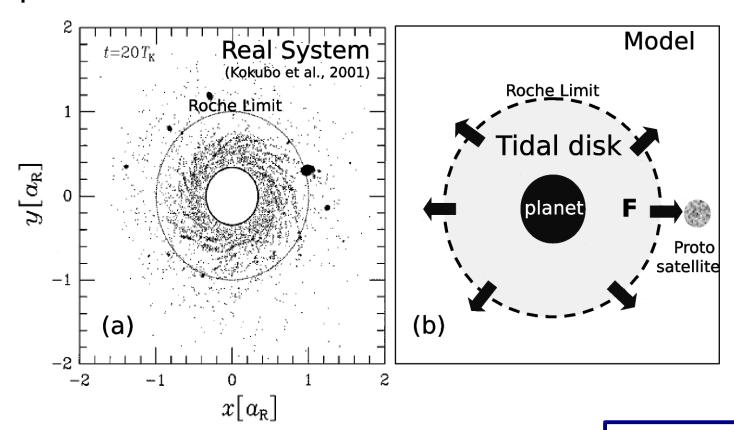


Crida & Charnoz 2012

## Spreading of a tidal disk

1D model.

Inside the Roche radius  $\mathbf{r}_{R}$ , there is a « tidal disk », that spreads with a mass flux  $\mathbf{F}$ .



#### **Notations**

Be  $T_R$  the orbital period at  $r_R$ , and

 $T_{disk} = M_{disk} / FT_{R}$ , the normalized life-time of the disk.

The disk spreads with a viscous time  $t_v = r_R^2/v$ .

Using Daisaka et al. (2001)'s prescription for v, we find  $T_{disk} = t_v / T_R = 0.0425 \ D^{-2}$  where  $D = M_{disk} / M_p$ , and  $F = 23 \ D^3 \ M_n / T_R$ .

## **Continuous regime**

Say 1 satellite forms. Its mass is : M = F t

$$M = F t \qquad (1)$$

It feels a torque from the tidal disk :  $\Gamma = \frac{8}{27} \left( \frac{M}{M_p} \right)^2 \Sigma r^4 \Omega^2 \Delta^{-3}$ where  $\Delta = (r-r_{\rm p})/r_{\rm p}$  (Lin & Papaloizou 1979).

→ Migration rate :

where  $q = M / M_p$ .

$$\frac{d\Delta}{dt} = \frac{32}{27} q D T_R^{-1} \Delta^{-3}$$
 (2)

Solution of (1) & (2) :

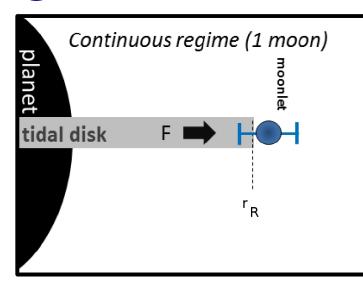
$$q = \left(\frac{\sqrt{3}}{2}\right)^3 T_{disk}^{-1/2} \Delta^2 \tag{3}$$

We call this the *continuous regime* .

# Continuous regime

This holds as long as the satellite captures immediately what comes through  $r_{\rm R}$ .

That is, as long as  $(r-r_R) < 2 r_{Hill}$ , or  $\Delta < 2 (q/3)^{1/3}$ .



Input into Eq.(3), this gives a condition of validity for the continuous regime :

$$\Delta < \Delta_c = \sqrt{\frac{3}{T_{disk}}} = -8.4 \text{ D}$$

$$q < q_c = \frac{3^{5/2}}{2^3} T_{disk}^{-3/2} = -222 \text{ D}^3$$

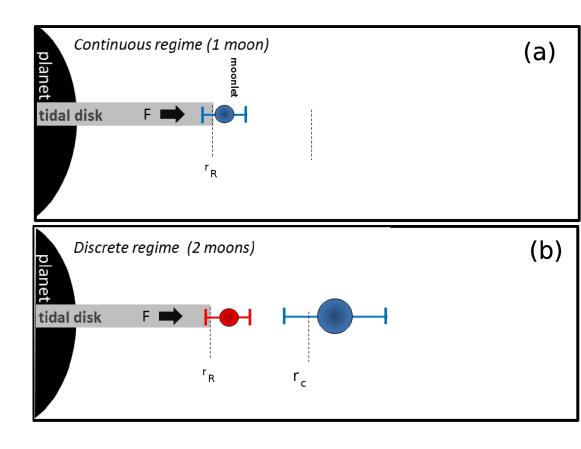
Duration of the continuous regime: 10  $T_R$ 

## Discrete regime

When the satellite is beyond  $\Delta_c$  (or  $q_c$ ), the material flowing through  $r_R$  forms a new satellite at  $r_R$ .

This new satellite is immediately accreted by the first one.

And so on...



The first satellite still grows as M=Ft, but by steps : discrete regime.

## Discrete regime

This holds as long as  $\Delta < \Delta_c + 2(q/3)^{1/3}$ .

It gives the condition:

$$\Delta < \Delta_d = \text{ r. If } \Delta_c = \text{~~26 D}$$
 
$$q < q_d = \text{~~1} q_c = \text{~~2200 D}^3$$

$$q < q_d = 9.9 q_c = -2200 D^3$$

The duration of the discrete regime is  $\sim 100 \text{ T}_{\scriptscriptstyle D}$ .

## Discrete regime

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It gives the condition:

$$\Delta < \Delta_d = 3.14 \Delta_c = \sim 26 D$$

$$q < q_d = 9.9 q_c$$
 = ~2200 D<sup>3</sup>

The duration of the discrete regime is  $\sim 100 T_{_{\rm R}}$ .

#### **Applications:**

- 1) Earth's Moon forming disk :  $q_d$  = mass of the Moon!
- 2) Charon never left the continuous regime.
- 3) Saturn's rings :  $q_d = \sim 10^{-18}$  .

Satellites of mass  $q_d$  are produced at  $\Delta_d$  every  $q_d/F$ .

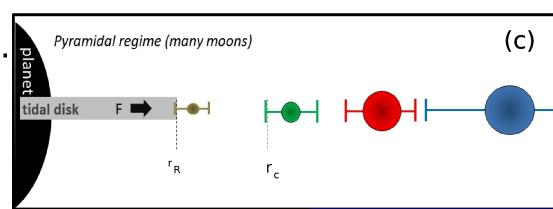
Then, many satellites of constant mass migrate outwards, at decreasing rates. They approach each other.

If their distance decreases below 2 mutual Hill radii, they merge.

This leads to the formation of satellites of masses  $2q_d$ , every  $2q_d$ /F. They migrate away and merge further...

And so on, hierachicaly...

We call this the pyramidal regime.



- Using Eq.(2), we show that in the pyramidal regime, while the mass is doubled,  $\Delta$  is multiplied by  $2^{5/9}$ .

Thus,  $q \propto \Delta^{9/5}$ 

In addition, the number density of satellites should be proportionnal to  $1/\Delta$ , explaining the pile-up.

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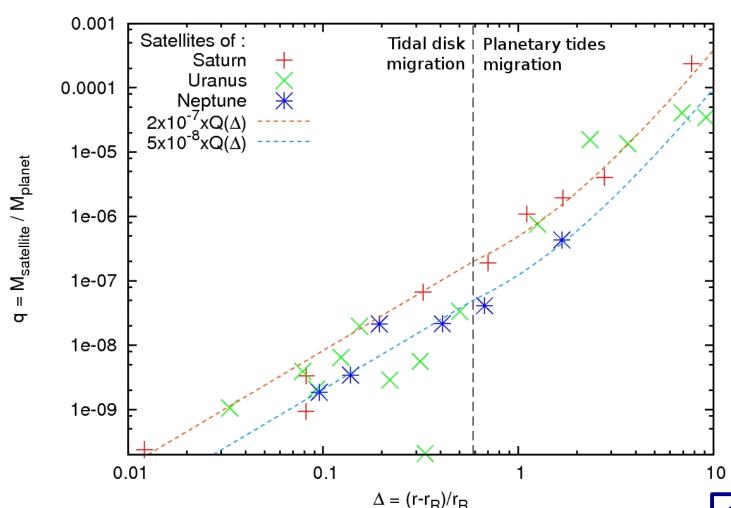
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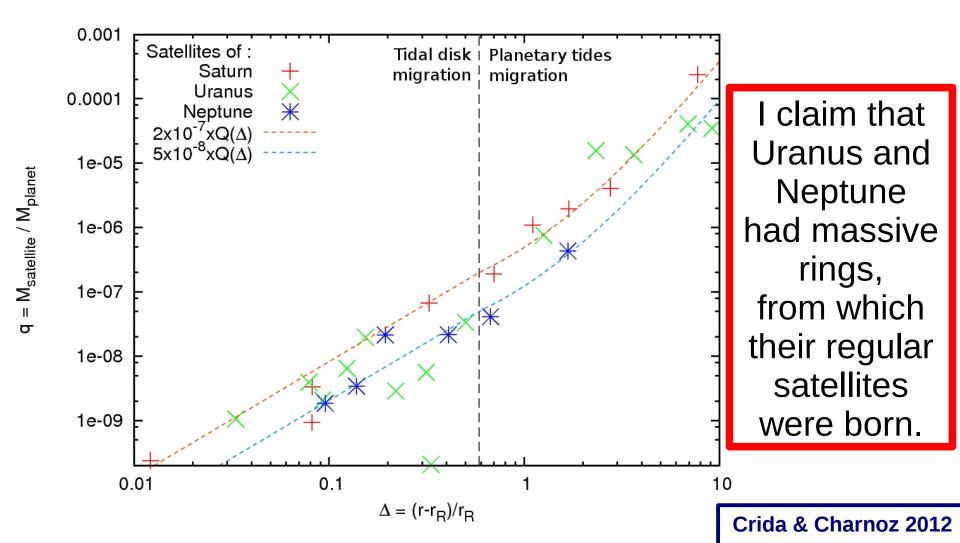
- Beyond the 2:1 Lindblad resonance with  $r_{p}$  ( $\Delta$ =0.58), Eq.(2) doesn't apply. Migration is driven by planetary tides:

$$\frac{dr}{dt} = \frac{3 k_{2p} M \sqrt{G} R_p^5}{Q_p \sqrt{M_p} r^{11/2}} \qquad (4)$$
 Using Eq.(4), we find q  $\alpha$  r  $^{3.8}$ 

The result spectacularly matches the distribution of the Saturnian, Uranian, and Neptunian systems!



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## **Summary**

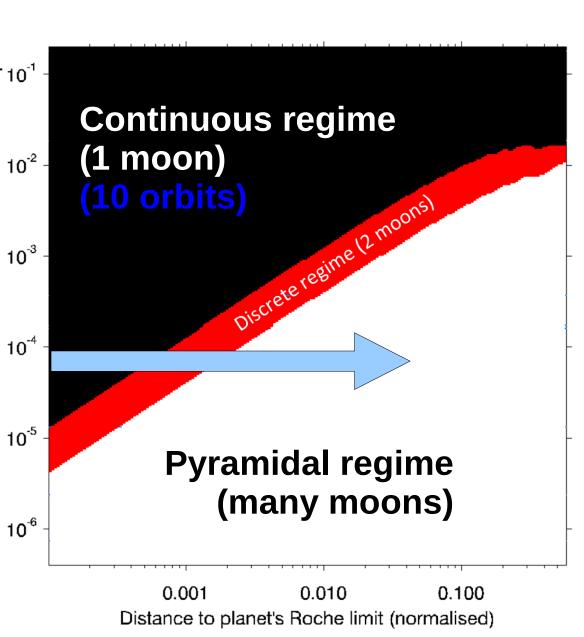
Disk-to-planet mass ratio

1) Continuous regime:  $_{10^{-1}}$  1 moon grows  $q \alpha \Delta^2$  until  $\Delta_c$  or  $q_c$ .

2) Discrete regime: 2 moons, growth by steps until  $\Delta_d$  or  $q_d$ .

**3) Pyramidal regime:** Many moons in the system.

q  $\alpha \, \Delta^{9/5}$  or  $r^{3.8}$  .

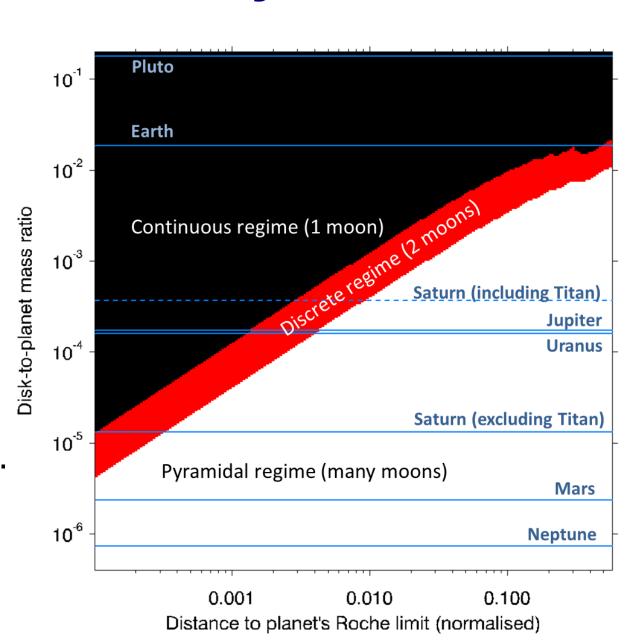


## **Summary**

Take  $M_{disk} = 1.5 x$ the mass of the present satellite system.

Giant planets must be dominated by the pyramidal regime,

while we expect the Earth and Pluto to have 1 large satellite.



#### **Conclusion & Discussion**

The spreading of a tidal disk beyond the Roche radius

- explains the mass-distance distribution of the regular satellites of the giant planets (observational signature of this process)
- unifies terrestrial and giant planets in the same paradigm.
- most Solar System regular satellites formed this way.

- \* Jupiter doesn't fit in this picture: probably formed in a circum-planetary disk (e.g. Canup & Ward 2002, 2006; Sasaki et al 2010)
- Titan fits very well in this picture, though its « tidal age » is too large... Coincidence ?



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