

INTERSTELLAR BUBBLES IN TWO YOUNG H II REGIONS

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ABSTRACT

Massive stars are expected to produce wind-blown bubbles in the interstellar medium; however, ring nebulae, suggesting the existence of bubbles, are rarely seen around main-sequence O stars. To search for wind-blown bubbles around main-sequence O stars, we have obtained high-resolution *Hubble Space Telescope* WFPC2 images and high-dispersion echelle spectra of two pristine H II regions, N11B and N180B, in the Large Magellanic Cloud. These H II regions are ionized by OB associations that still contain O3 stars, suggesting that the H II regions are young and have not hosted any supernova explosions. Our observations show that wind-blown bubbles in these H II regions can be detected kinematically, but not morphologically, because their expansion velocities are comparable to or only slightly higher than the isothermal sound velocity in the H II regions. Bubbles are detected around concentrations of massive stars, individual O stars, and even an evolved red supergiant (a fossil bubble). Comparisons between the observed bubble dynamics and model predictions show a large discrepancy (1–2 orders of magnitude) between the stellar wind luminosity derived from bubble observations and models and that derived from observations of stellar winds. The number and distribution of bubbles in N11B differ from those in N180B, which can be explained by the difference in the richness of stellar content between these two H II regions. Most of the bubbles observed in N11B and N180B show a blister structure, indicating that the stars were formed on the surfaces of dense clouds. Numerous small dust clouds, similar to Bok globules or elephant trunks, are detected in these H II regions, and at least one of them hosts on-going star formation.

Key words: H II regions — ISM: bubbles — ISM: individual (N11B, N180B) —
 ISM: kinematics and dynamics — Magellanic Clouds

1. INTRODUCTION

The fast wind of a massive star can sweep the ambient medium into a shell and form a wind-blown bubble (Weaver et al. 1977). If the ambient density is sufficiently high, the swept-up shell should appear as a “ring nebula” in H α images. Ring nebulae are frequently seen around evolved stars with copious stellar winds, such as Wolf-Rayet (WR) stars and luminous blue variables (LBVs); however, ring nebulae are rarely seen around unevolved massive stars, such as main-sequence O stars (Chu 1991). In fact, no main-sequence O stars in the Galaxy possess well-defined optical ring nebulae.

The paucity of optical ring nebulae around unevolved massive stars is puzzling, as early-type O dwarfs may have stellar winds with terminal velocities of $\sim 2000 \text{ km s}^{-1}$ and mass-loss rates of a few times $10^{-6} M_{\odot} \text{ yr}^{-1}$ (Prinja, Barlow, & Howarth 1990; de Jager, Nieuwenhuijzen, & van der Hucht 1988). Three possible explanations exist: (1) no careful, systematic searches for ring nebulae around main-sequence massive stars have been conducted, (2) main-

sequence massive stars preferentially reside in low-density media, and (3) the microstructure of the interstellar medium prohibits the formation or the visibility of wind-blown bubbles (e.g., McKee, van Buren, & Lazareff 1984).

To search for bubbles blown by unevolved massive stars, we have chosen to observe two young OB associations in dense H II regions in the Large Magellanic Cloud (LMC). OB associations are selected so that a large number of stars can be imaged in the same frame. The “young age” criterion is used to avoid confusing supernova remnants (SNRs) and superbubbles, which are formed by evolved stars. The dense H II region environment is required to ensure a reasonably dense medium around the massive stars. The two OB associations we have selected are LH10 and LH117 (Lucke & Hodge 1970) in the H II regions N11B⁷ and N180B (Henize 1956), respectively. Both OB associations contain O3 stars (Parker et al. 1992, hereafter PGMW; Massey et al. 1989, hereafter MGSD). The presence of O3 stars indicates that even the most massive stars have remained intact, thus suggesting a young age and enhancing the likelihood of a real absence of previous supernovae and justifying the “pristine” state of the H II region.

The observations we have carried out for the H II regions N11B and N180B include high-resolution images obtained with the *Hubble Space Telescope* Wide Field Planetary Camera 2 (*HST* WFPC2) and high-dispersion long-slit

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⁷ N11B is a compact H II region in the N11 H II complex. While N11 contains a known SNR, N11L, it is at 180 pc away and is totally unrelated to N11B.

spectra obtained with an echelle spectrograph on the 4 m telescope at Cerro Tololo Inter-American Observatory (CTIO). Despite the high spatial resolution of the *HST* observations, few ring nebulae are found around main-sequence massive stars in LH10 and LH117.

In this paper we report the analysis of *HST* and ground-based observations of the H II regions N11B and N180B in the LMC, and we use these images and spectra to investigate the issue of ring nebulae around main-sequence O stars. The observations are described in § 2, the analysis and results are reported in § 3, and the conclusions are discussed in § 4.

2. OBSERVATIONS

The data sets used in this study include (1) images taken with the *HST* WFPC2 and (2) high-dispersion echelle spectra taken with CTIO 4 m telescope. The high-dispersion spectra are useful to diagnose expanding shells that are not morphologically identifiable.

2.1. *HST* WFPC2 Images

HST WFPC2 images of N11B and N180B were taken on 1999 May 12 and 1998 April 29, respectively, for the Cycle 6 program 6698. For each H II region, the observations were made through the F502N filter for 2×600 s and the F656N filter for 2×500 s. The F502N filter, centered at 5012.2 Å with a FWHM of 26.8 Å, includes only one nebular line, [O III] λ 5007. The F656N filter, centered at 6563.7 Å with a FWHM of 21.4 Å, may include the H α line and the neighboring [N II] lines. At the LMC's radial velocity, ~ 300 km s $^{-1}$, the [N II] λ 6583 line is redshifted further away from the red edge of F656N's bandpass, but the [N II] λ 6548 line is redshifted further into F656N's bandpass. The filter transmission for the redshifted [N II] λ 6548 is 95% that at the redshifted H α . Using the [N II]/H α intensity ratio measured from our echelle spectra (see § 2.2), we find that the [N II] λ 6548 line emission contributes to 1%–2% of the F656N flux of N11B and 2%–4% of that of N180B. This [N II] contamination has been corrected in our H α flux measurements.

The calibrated WFPC2 images were produced by the standard *HST* pipeline. We process them further with IRAF and STSDAS routines. The images taken with the same filter were combined to remove cosmic rays and to produce a total-exposure map. The combined [O III] and H α images were then corrected for the intensity- and position-dependent charge transfer efficiency by applying a linear ramp with a correction factor chosen according to the average counts of the sky background (Holtzman et al. 1995). Following the procedures for narrowband WFPC2 photometry,⁸ images were first divided by their total exposure time and then multiplied by the PHOTFLAM parameter of the header in order to get flux densities. To obtain fluxes, we multiplied the flux densities by the filter rectangular width calculated with SYNPHOT: 35.8 Å for the F502N filter and 28.3 Å for the F656N filter.

Figures 1a and 1b show the large-scale environments of N11B and N180B with the WFPC2 fields marked. Figures

1c and 1d show the WFPC2 H α images of N11B and N180B marked with the identifications and spectral types of the most relevant stars in their OB associations (LH 10 in N11B, PGMW; LH 117 in N180B, MGSD). To better show the morphological features, we present the WFPC2 H α images of the two nebulae individually in Figures 2 and 3. To show the excitation variations, the [O III] images and [O III]/H α ratio map of N11B and N180B are shown in Figures 4 and 5, respectively.

2.2. CTIO 4 Meter Echelle Spectra

Echelle spectra were obtained with the CTIO 4 m telescope in 2000 January and December. The spectrograph was used with a 79 line mm $^{-1}$ echelle grating and the long-focus red camera. We observed a single order centered on the H α line by inserting a post-slit H α interference filter and replacing the cross-dispersing grating with a flat mirror. A Tek 2048 \times 2048 CCD with 24 μ m pixel $^{-1}$ was used to record the images. This provided a spectral sampling of 0.08 Å pixel $^{-1}$ and a spatial sampling of 0".26 pixel $^{-1}$. The wavelength coverage, limited by the H α filter and the echelle order, was 125 Å; the spatial coverage, limited by the optics, is $\sim 200''$. The angular resolution, determined by the FWHM of the seeing, was approximately 1". A slit width of 1".65 was used. The resultant spectral resolution, measured from the Th-Ar lamp lines, is about 13 km s $^{-1}$ FWHM.

The journal of observations is given in Table 1. When multiple exposures existed for a given position, they were combined, taking into account their exposure times.⁹ On the whole, nine slit positions were observed in N11B and four in N180B. They were all oriented either north-south or east-west; their exact positions were indicated by the arrows in Figures 1c and 1d. The identifications of lines detected in the echelle observations are shown in Figure 6.

⁹ Two observations east-west of PGMW 3168 were taken. Since they were not centered at the same position, we did not combine them and instead analyzed them individually.

TABLE 1
JOURNAL OF ECHELLE OBSERVATIONS

Central Star	Slit Orient.	Exp. Time (s)	Obs. Date
N11B:			
PGMW 3204.....	East-west	600	2000 Jan 23
PGMW 3224.....	East-west	600	2000 Jan 23
PGMW 3168.....	East-west	2×600^a	2000 Jan 23
PGMW 3058.....	North-south	2×300	2000 Jan 22
PGMW 3053.....	North-south	2×300	2000 Jan 22
PGMW 3157.....	North-south	$2 \times 300 + 10$	2000 Jan 22
PGMW 3223.....	East-west	600	2000 Jan 23
PGMW 3120.....	North-south	2×300	2000 Jan 22
PGMW 3120.....	East-west	600	2000 Jan 23
N180B:			
MGSD 140.....	East-west	600	2000 Jan 23
MGSD 118.....	North-south	600	2000 Dec 10
MGSD 214.....	North-south	600	2000 Dec 10
MGSD 214.....	East-west	600	2000 Jan 23

⁸ Available at http://www.stsci.edu/instruments/wfpc2/Wfpc2_faq/wfpc2_nrw_phot_faq.html.

^a Two observations of PGMW 3168 were taken, but they were not centered at the same position.

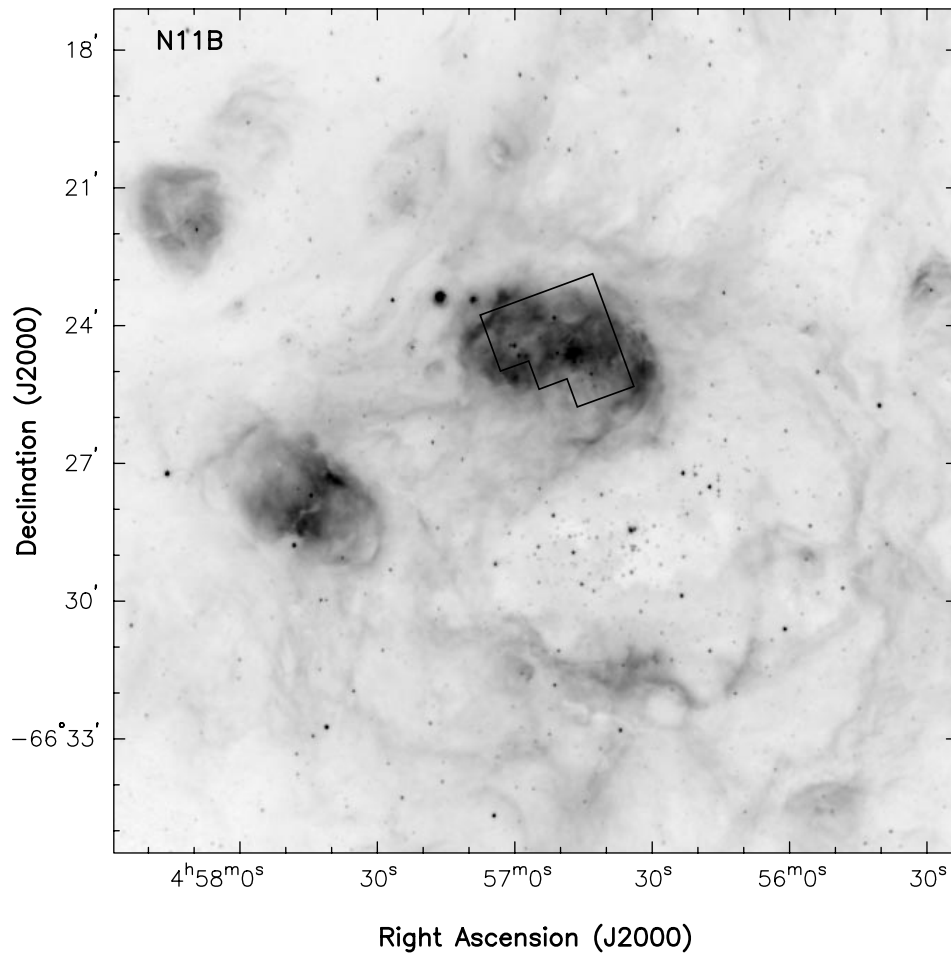


FIG. 1a

FIG. 1.—(a) Curtis Schmidt H α CCD image of N11B with the WFPC2 field outlined. (b) Same as (a) for N180B. (c) *HST* WFPC2 H α image of N11B. Stars cataloged by PGMW are marked with their catalog number and spectral classification. Slit positions of the echelle spectra are indicated by arrows. (d) Same as (c) for N180B. The star numbers and spectral types are taken from MGSD. Images presented in this article, except for Fig. 7, are negative images.

The spectral lines include nebular H α and [N II] lines and telluric OH lines (Osterbrock et al. 1996). The rest wavelengths of the nebular lines we have adopted are H α λ 6562.7885 Å and [N II] λ 6548.0800 Å and λ 6583.4540 Å (Spyromilio 1995).

3. ANALYSIS AND RESULTS

For each nebula, we first describe the overall H α morphology and surface brightness variations and derive the rms electron density for various regions. We then identify morphological features that might have been produced by stellar winds interacting with the ambient interstellar medium, as seen in the WFPC2 H α images. We further use the [O III]/H α ratio map to determine whether a morphological feature has an anomalous [O III]/H α ratio. An elevated [O III]/H α ratio is indicative of a higher temperature or a shock excitation, while a lower [O III]/H α ratio suggests a lower excitation or ionization. Next we describe the kinematic features detected in the echelle spectra, fit Gaussian components to the line profiles, and use the resultant velocities to identify expanding structures and to determine expansion velocities. Finally, we combine the morphological and kinematic information and present an integrated view of the structure of the H II region.

3.1. N11B

3.1.1. Surface Brightness and Density

The H II region N11B is a young star-forming region dotted with numerous bright-rimmed dusty features. The OB association LH10 has a high concentration of stars at the southwestern part of N11B, centered near $4^{\text{h}}56^{\text{m}}44^{\text{s}}$, $-66^{\circ}25'00''$ (see Fig. 1c). To the north of this concentration is the brightest H α emission region mixed with dusty features and connected with several long filaments. The other members of LH10 are loosely distributed within N11B, together with small (~ 1 pc in size) dust clouds that are similar to Bok globules (Garnett et al. 1999) or elephant trunks (Hester et al. 1996). The two most prominent small dust clouds are situated close to luminous stars, so their surfaces are ionized and emit strongly in H α : the kiwi-shaped dust cloud to the west of PGMW 3204 and PGMW 3209 and the Y-shaped dust cloud to the west of PGMW 3223 (see Fig. 1c). Interestingly, the tip of the Y-shaped dust cloud harbors a bright star, and our echelle observations (§ 3.1.3) show that this bright star has a strong H α emission line. Several smaller dust clouds are present, but at larger distances from luminous O stars, so their H α surface brightnesses are not as high as those close to luminous O stars.

To quantitatively analyze these images, we have mea-

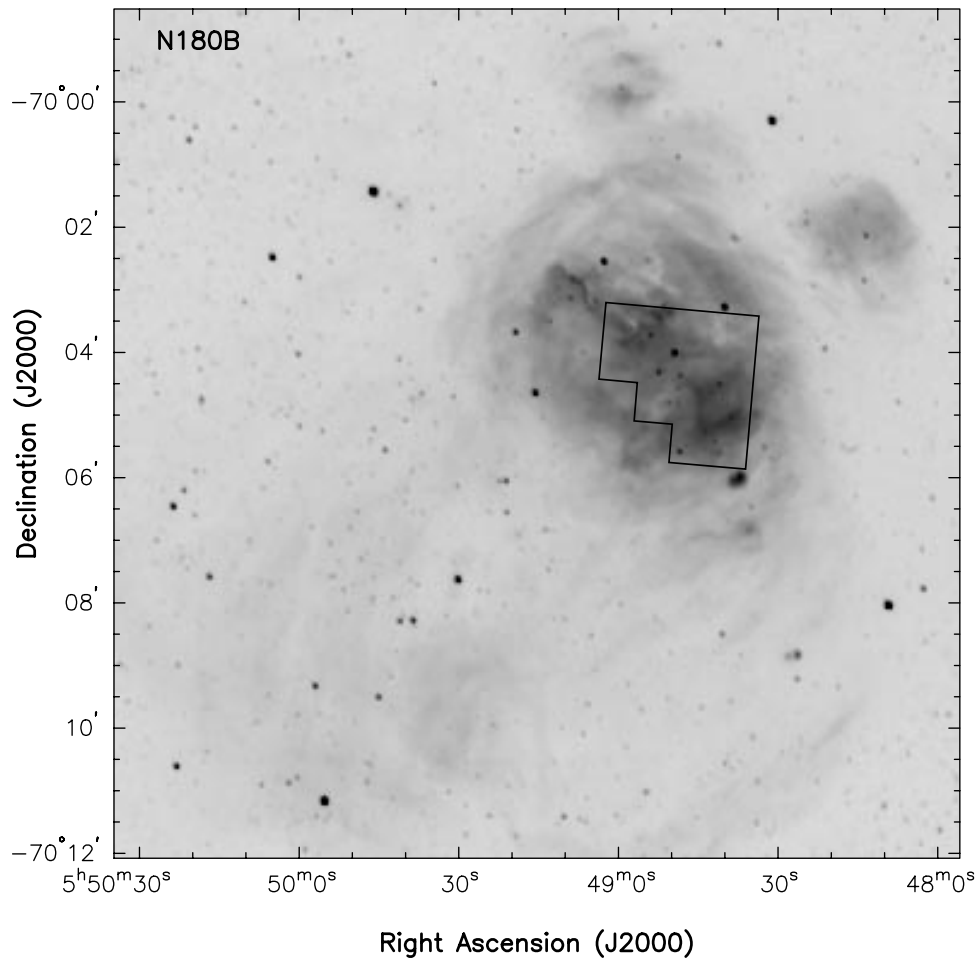


FIG. 1b

sured the $H\alpha$ surface brightness in several regions of interest. We then compute the emission measure (EM) using $EM = 2.41 \times 10^3 T^{0.92} S(H\alpha) \text{ pc cm}^{-6}$, where T is the electron temperature and $S(H\alpha)$ is the intrinsic surface brightness in units of ergs per square centimeter per second per steradian (Peimbert, Rayo, & Torres-Peimbert 1975). Adopting an electron temperature of 10^4 K and assuming that the depth of emitting material along the line of sight is similar to the observed size of a morphological feature, we may further estimate the rms electron density of the emitting material. If the feature is not spherical, the observed width (the shorter dimension) and length (the longer dimension) are used as the path length of emission to determine the upper and lower limits of the rms density, respectively.

We assume that the density structure of N11B consists of local density enhancements superposed on a global, uniform component. For the global component, we have measured the surface brightness of a featureless region at $9''$ east and $1'4''$ north of PGMW 3168, and derived $S(H\alpha) = 9.6 \times 10^{-4} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}$, corresponding to an EM_{global} of $1.1 \times 10^3 \text{ cm}^{-6} \text{ pc}$, taking into account an interstellar absorption of $E(B-V)=0.17$ mag (PGMW). The size of N11B, measured from the $H\alpha$ image in Figure 1a, is $257'' \times 140''$, or $64 \text{ pc} \times 35 \text{ pc}$ for a distance of 50 kpc. The rms electron density of the global component in N11B is thus 10 to 18 cm^{-3} in N11B. We have measured the

density enhancements in several arcs and the bright ionized surface of the kiwi-shaped dust cloud. The intrinsic surface brightness, excess emission measure ($EM - EM_{\text{global}}$), size of the feature, and the range of rms density of these features are listed in Table 2. It can be seen that the moderately bright arcs and filaments have rms densities of $50\text{--}150 \text{ cm}^{-3}$, while the bright ionized surfaces of dust clouds have rms densities several times higher.

3.1.2. Morphology and Excitation

Whereas there are a large number of O stars in N11B, only two of them (PGMW 3120 and PGMW 3160) are surrounded by filamentary features that are commonly identified as ring nebulae. The star PGMW 3120, classified as O5.5 V((f*)) by PGMW, is surrounded in the northeast by a small arclike structure at $\sim 5''$ from the star. The arc does not stand out in the $[O\text{ III}]/H\alpha$ ratio map, suggesting a similar excitation to the background H II region or a lack of strong shock excitation. The star PGMW 3160, a K I star, is in a faint void bordered in the southwest by a bright filament $\sim 20''$ long. This filament has a lower $[O\text{ III}]/H\alpha$ ratio than the background H II region; however, the filament does not appear to be associated with an obvious dusty feature.

A few stars are surrounded by faint filaments, e.g., PGMW 3239 (B2 V), PGMW 3157 (BC1 Ia), and PGMW 3102 (O7 V). Some stars, e.g., PGMW 3089 (O8 V), are

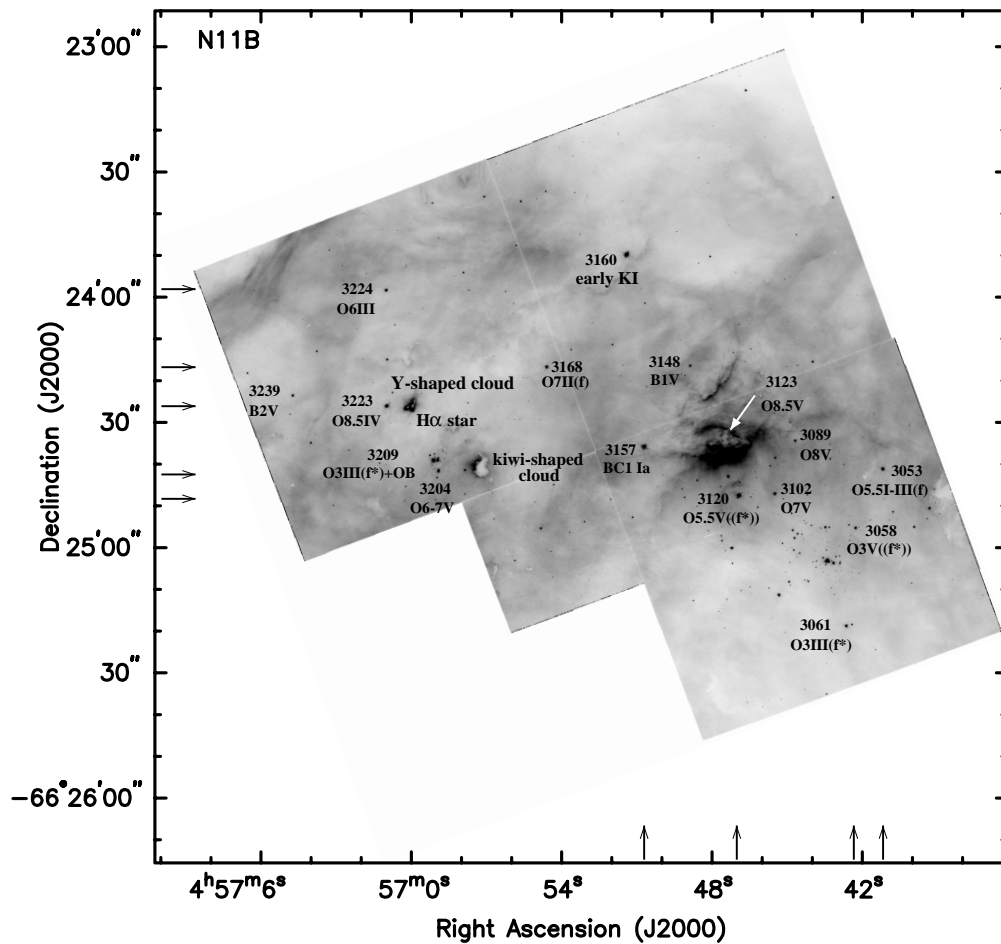


FIG. 1c

inside regions with suppressed surface brightness, indicating lower densities around the star, as in a cavity. None of these features stand out in the $[\text{O III}]/\text{H}\alpha$ ratio map; therefore, from their excitation alone, they are indistinguishable from the background H II region, indicating that they are not shock excited.

While no sharp, high- $[\text{O III}]/\text{H}\alpha$ features are detected, many prominent low- $[\text{O III}]/\text{H}\alpha$ features are present, and

most of them are associated with dusty features. The correlation between low $[\text{O III}]/\text{H}\alpha$ ratios and dusty features can be illustrated by the analysis of an isolated dust cloud. We use the kiwi-shaped dust cloud as an example. Figure 7 shows that the low- $[\text{O III}]/\text{H}\alpha$ region delineates the surface of the dust cloud. The $\text{H}\alpha$ and $[\text{O III}]$ surface brightness profiles along a north-south cut across the kiwi-shaped dust cloud show that the $\text{H}\alpha$ emission and $[\text{O III}]$ emission on

TABLE 2
DENSITIES OF SELECTED REGIONS IN N11B AND N180B

Region ^a	$S(\text{H}\alpha)$ ($10^{-3} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}$)	$\text{EM} - \text{EM}_{\text{global}}$ ($10^4 \text{ cm}^{-6} \text{ pc}$)	Size (arcsec \times arcsec)	rms Density (cm^{-3})
N11B:				
Near PGMW 3168 (background)	0.955 ± 0.016	$(1.101 \pm 0.019)^b$	257×140	13–18
Arc south of PGMW 3160	1.823 ± 0.020	1.001 ± 0.030	17.8×1.5	50–160
Arc west of PGMW 3148	3.280 ± 0.028	2.683 ± 0.037	11.5×2.2	100–220
Arc north of PGMW 3053	2.288 ± 0.025	1.538 ± 0.035	7.8×2.8	90–150
Bright region of the kiwi cloud	5.096 ± 0.127	4.776 ± 0.147	2.6×0.8	270–500
N180B:				
Near MGSD 119 (background)	0.424 ± 0.003	$(0.489 \pm 0.004)^b$	225×225	9.5
South of MGSD 214	0.297 ± 0.002	$(0.343 \pm 0.002)^c$	225×225	8
Arc north of MGSD 218	0.696 ± 0.007	0.314 ± 0.009	5.4×1.3	50–100
V-shape zone between MGSD 187 and MGSD 197	0.764 ± 0.002	0.392 ± 0.005	37.0×7.0	20–50

^a The different regions can be identified in Figs. 1c and 1d.

^b This is the background reference, $\text{EM}_{\text{global}}$, and the size of this feature is that of the nebula.

^c This EM is not background subtracted.

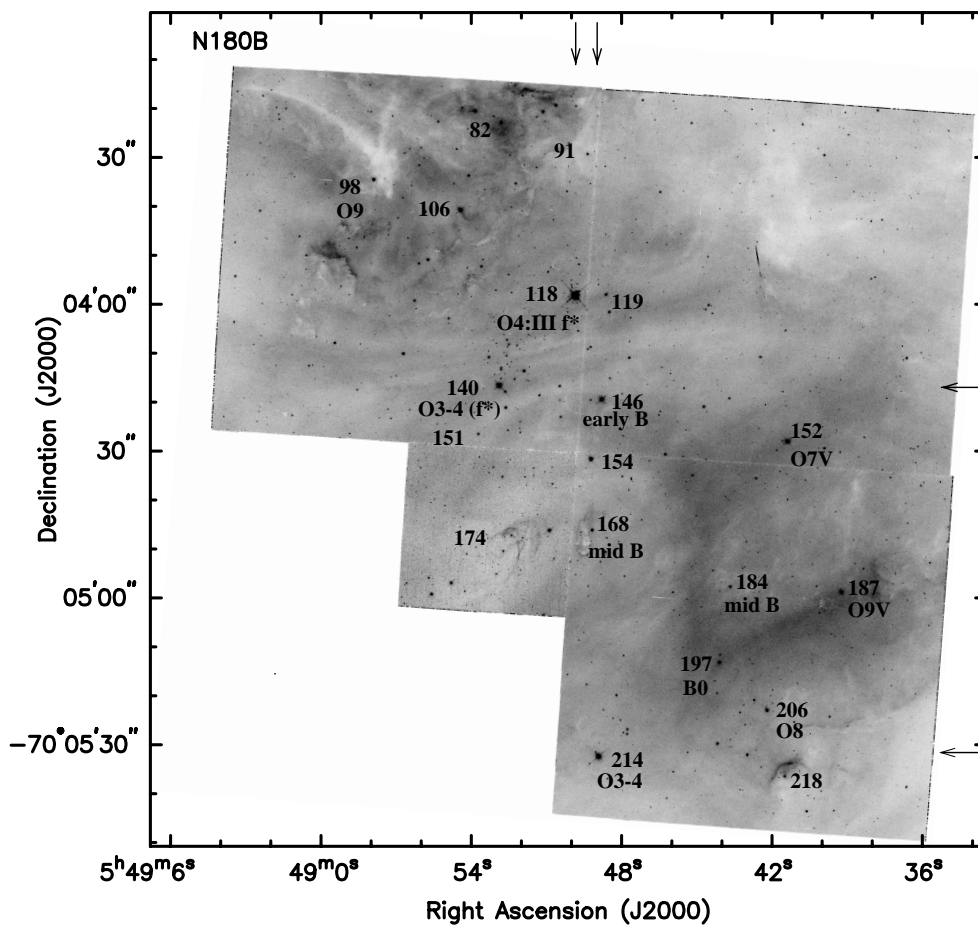


FIG. 1d

the cloud surface peak at the same location, but the $[\text{O III}]$ emission drops off faster into the dust cloud than the $\text{H}\alpha$ emission, causing a dip in the $[\text{O III}]/\text{H}\alpha$ ratio. Therefore, the low $[\text{O III}]/\text{H}\alpha$ ratios on the surface of a dust cloud are caused by a lower ionization and excitation. This kiwi-shaped dust cloud and the Y-shaped dust cloud are both strongly ionized on the sides facing the O3 III(f*) star PGMW 3209, the earliest and most luminous O star in the vicinity. These dust clouds are most likely illuminated and ionized by this O star.

3.1.3. Kinematic Properties

To identify kinematic features caused by wind-ISM interactions, we examine the velocity profiles of both $\text{H}\alpha$ and $[\text{N II}]$ $\lambda 6584$ lines. The $[\text{N II}]$ line shows line splitting more easily, because its thermal width is smaller than that of $\text{H}\alpha$, but the $[\text{N II}]$ line is much weaker than the $\text{H}\alpha$ line and is usually noisier. Variations of the ionization condition ($[\text{N II}]/\text{H}\alpha$ ratio) among different velocity components also contribute to the differing appearances of the $\text{H}\alpha$ and $[\text{N II}]$ line profiles. These effects are illustrated in Figure 8, where two pairs of $\text{H}\alpha$ and $[\text{N II}]$ line profiles are shown.

We fit Gaussian components to both $\text{H}\alpha$ and $[\text{N II}]$ profiles; the resultant velocity components plotted along each slit are presented in Figure 9. To compare the kinematic and morphological properties, in Figure 10 we mark the regions with more than one velocity component along the slits, with the brighter components marked in thick solid

lines and the weaker components in thick dashed lines. All quoted velocities are heliocentric. Four echellograms of the $[\text{N II}]$ line with line splitting are shown in Figure 11.

A completely different view emerged when we examined the echelle spectra and used the kinematic information to identify features caused by wind-ISM interactions. The kinematic structure of N11B has been studied by Rosado et al. (1996) using imaging Fabry-Perot observations at an $18''$ resolution ($9'' \text{ pixel}^{-1}$). They reported six expanding regions (designated "a"–"f"). To study these expanding regions at a higher angular resolution, we have chosen our first six slit positions centered on stars in these regions. The other slit positions were selected to study morphologically identified features: the Y-shaped dust cloud and the small ring nebula around PGMW 3120. Below we review the kinematic features along individual slits. The spectral types of the reference stars and the names of the expanding regions of Rosado et al. (1996) are noted in parentheses.

1. East-west of PGMW 3204 (O6–7 V; Region a). The $\text{H}\alpha$ and $[\text{N II}]$ lines are significantly broader in the region near the O3 III(f*) star PGMW 3209. The $\text{H}\alpha$ and $[\text{N II}]$ lines can be fitted by two Gaussian components centered at 288 km s^{-1} and 310 km s^{-1} . The velocity plot shows a clear expanding shell structure, with an expansion velocity of $\sim 10 \text{ km s}^{-1}$. This expanding shell can be traced from $2''$ west to $54''$ east of PGMW 3204, corresponding to a linear size of 14 pc. Interestingly, the kiwi-shaped dust cloud is located right at the western edge of the expanding shell,

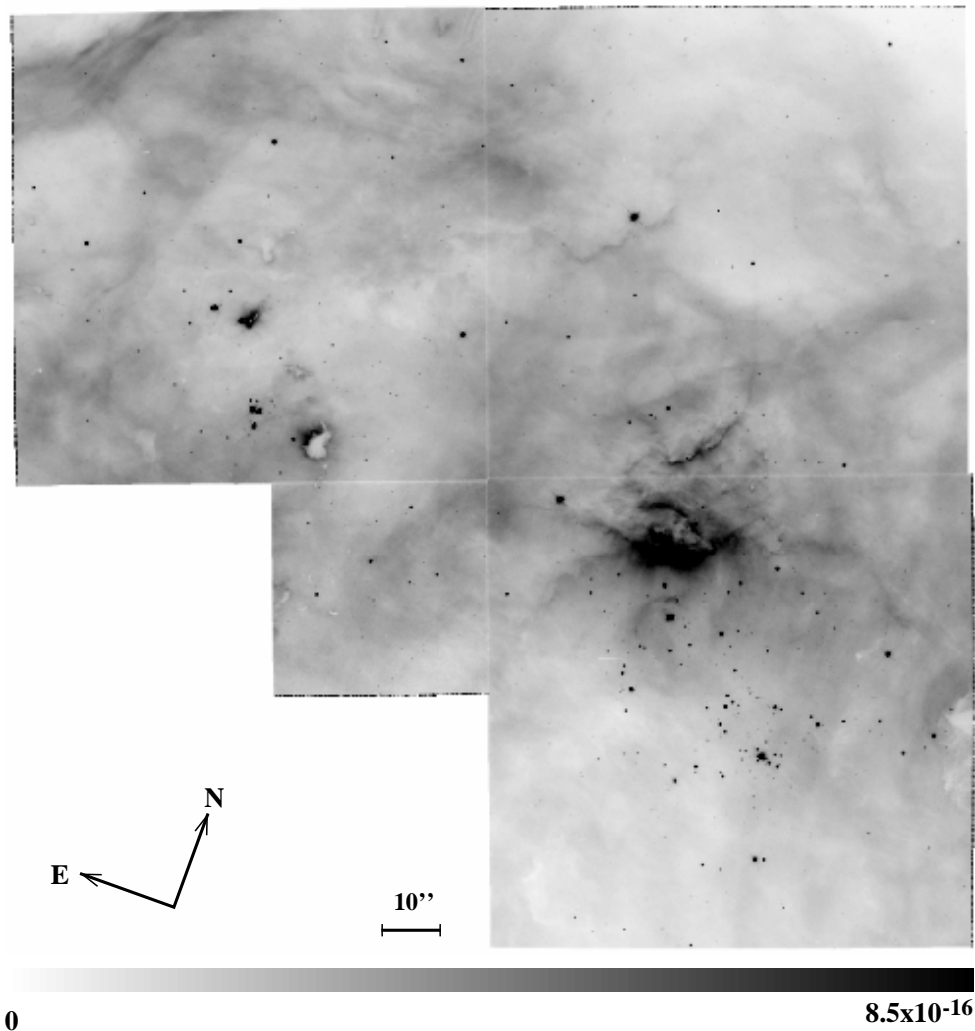


FIG. 2.—*HST* WFPC2 $H\alpha$ image of N11B. The bar below the image shows the gray scale for surface brightness in units of ergs per square centimeter per second per WFC pixel.

suggesting again that this dust cloud is physically close to PGMW 3204. Despite the clear expanding shell structure revealed in the velocity plot, there are no morphological features in the WFPC2 $H\alpha$ image that can be unambiguously identified as a ring nebula or an expanding shell. Outside this expanding shell, both $H\alpha$ and $[N II]$ lines can be fitted well by a single Gaussian component, centered at 298 km s^{-1} on the east side and 293 km s^{-1} on the west side of the shell. The western end of the slit reaches the bright H II region north of PGMW 3120. The velocity profiles in the bright H II region are more or less Gaussian, peaking at 297 km s^{-1} , but additional faint components are detected in the region bounded by the two arcs extending from the H II region to the east.

2. East-west of PGMW 3224 (O6 III; Region b). No clear line splitting is seen in the spectra along the slit, but an additional faint blue component sometimes creates a small asymmetry in the line profiles. From $63''$ east to $7''$ west of PGMW 3224, both lines display a faint blue component at 282 km s^{-1} , while the brighter component peaks at 303 km s^{-1} . Further east, the $H\alpha$ line shows only one peak at 299 km s^{-1} . Further west, the line profiles peak at 299 km s^{-1} but become asymmetric again at $38''$ to $60''$ west of PGMW

3224, in the vicinity of PGMW 3160 (early K I). At the western end of this slit position, the $H\alpha$ line shows an additional component at redshifted velocities; this component is too faint to be detected in the $[N II]$ line.

3. East-west of PGMW 3168 (O7 II(f); Region c). Along this slit position, no clear line splitting is seen, but the peak velocities show a large scatter and the line profiles are often asymmetric, indicating multiple velocity components. At $133''$ to $50''$ east of PGMW 3168, both lines show Gaussian profiles centered on 298 km s^{-1} . At $50''$ to $23''$ east of the star, the $H\alpha$ line is broad and shows large velocity variation, while the $[N II]$ line profiles are noticeably asymmetric and require two velocity components in the spectral fits. The spectral fits to the $[N II]$ line show an expanding shell with an expansion velocity of 12 km s^{-1} and a diameter of $\sim 40''$ (or 10 pc). This shell might be associated with the uncataloged star between PGMW 3223 and PGMW 3224. Near PGMW 3168 and to its west, the $[N II]$ line also presents a main component (at 298 km s^{-1}) and a fainter blue component (at 283 km s^{-1}), although the $H\alpha$ line shows Gaussian profiles centered at 295 km s^{-1} . Further west, near the bright H II region, the $H\alpha$ line also shows asymmetric profiles, indicating velocity components at 290 and 307 km s^{-1} .

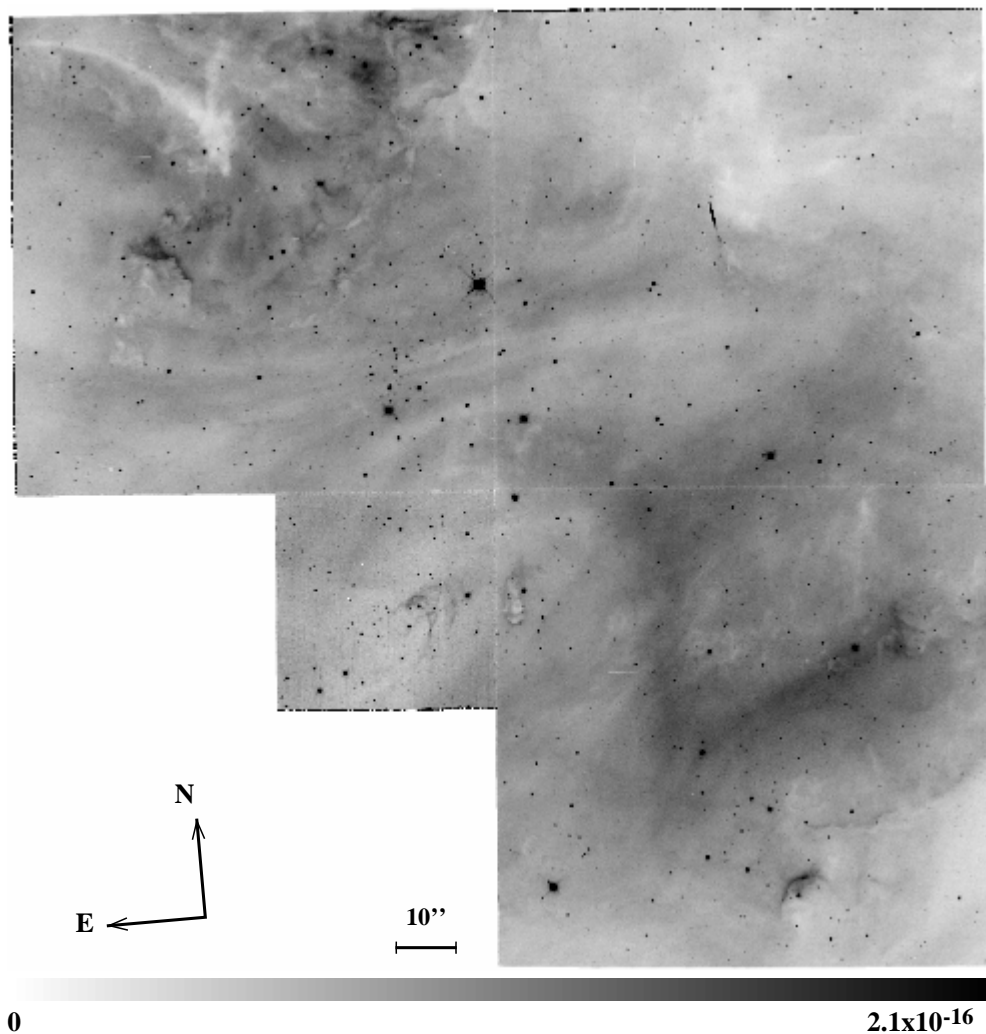


FIG. 3.—Same as Fig. 2, for N180B

Still further west (more than $67''$ west of the star), both lines seem split, but the red component is quite faint. This kinematic feature does not have any morphological counterparts in the WFPC2 images.

4. North-south of PGMW 3058 (O3 V((f*)); Region d).¹⁰ From $8''$ north of PGMW 3058 northward, both H α and [N II] lines exhibit clear line splitting. It is interesting to note that a bright knot appears in the red component at $25''$ north of PGMW 3058, where the WFPC2 image shows a bright filament. (An example of such emission knot can be seen in the echelle image of PGMW 3053 slit in Fig. 11.) South of the star the H α line is clearly split, with extreme velocities reaching 283 and 325 km s^{-1} , indicating an expansion velocity of $\sim 20 \text{ km s}^{-1}$. At $13''$ to $40''$ south of the star, in the vicinity of PGMW 3061 (O3 III(f*)), an additional, high-velocity component at 357 km s^{-1} is detected. It is not clear whether this high-velocity feature is accelerated by the powerful stellar wind of PGMW 3061, because a similar high-velocity feature is detected along the

north-south slit at PGMW 3120, only at a declination $20''$ south of PGMW 3061 (see description below of the slit north-south of PGMW 3120).

5. North-south of PGMW 3053 (O5.5 I–III(f); Region e). This slit position is only $8''$ east of the slit along north-south of PGMW 3058, thus showing a very similar kinematic structure. The line splittings and high-velocity features detected along the slit centered on PGMW 3058 are all confirmed in the spectra along the slit centered on PGMW 3053. The bright emission knot in the red component that corresponds to a bright filament to the north of PGMW 3053 is clearly seen in Figure 11.

6. North-south of PGMW 3157 (BC1 Ia; Region f). To the north of the star, the line profiles are mostly Gaussian and centered at 297 km s^{-1} ; however, at $43''$ to $56''$ north of the star, an additional faint blue component appears at 284 km s^{-1} , and the main component becomes slightly redshifted to 301 km s^{-1} . This region, as marked in Figure 10, corresponds to the arc feature around PGMW 3160 (early K I). Similar velocity structure is seen in the vicinity of PGMW 3160 in the echelle slit east-west of PGMW 3224. From PGMW 3157 southward to $29''$ south, an additional red component is present at $\sim 310 \text{ km s}^{-1}$. This is the

¹⁰ Note that Rosado et al.'s (1996) region d is centered on PGMW 3061, about $22''$ south of PGMW 3058.

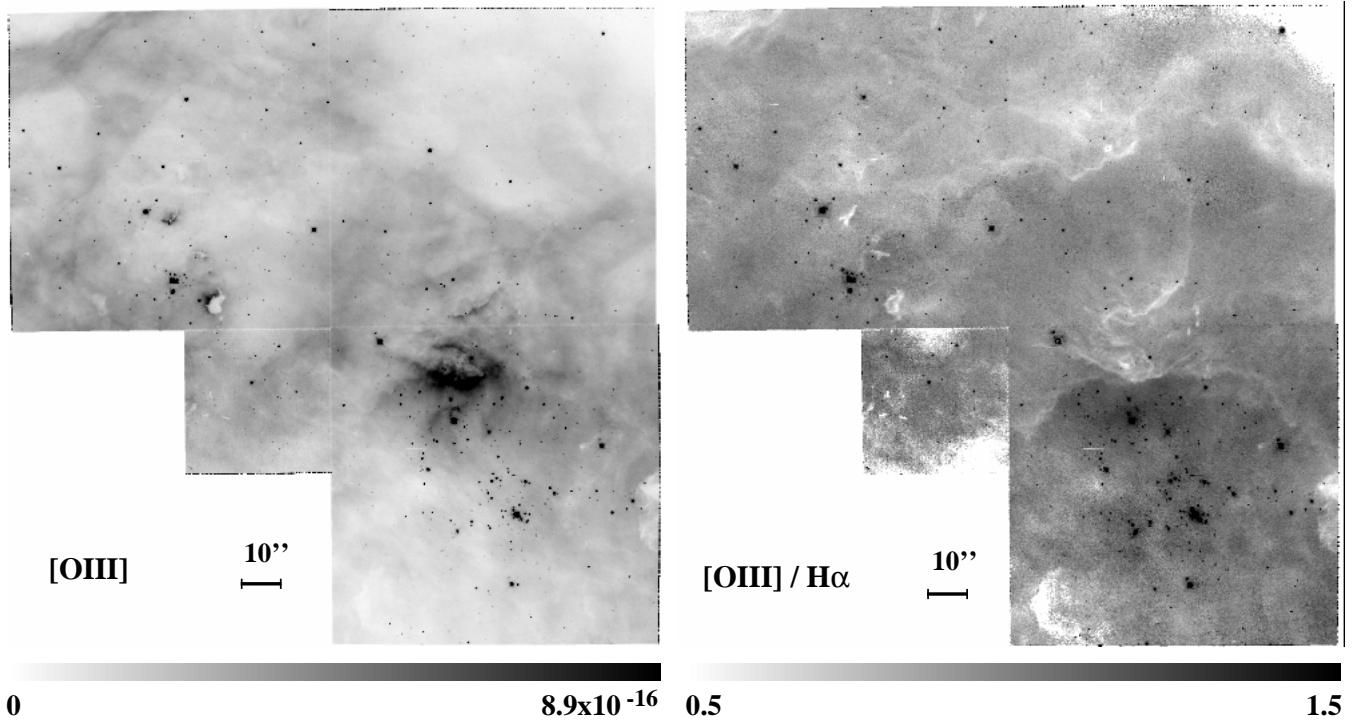


FIG. 4.—*HST* WFPC2 [O III] image (*left*) and [O III]/H α ratio map (*right*) of N11B. The bars below the images show the gray scales for surface brightness (in units of ergs per square centimeter per second per WFC pixel) and line ratio, respectively.

region to the south of a bright filament (see Fig. 10). A similar velocity structure in the vicinity is seen in the observation along the slit east-west of PGMW 3204.

7. East-west of PGMW 3223 (O8.5 IV). This slit position is selected to sample the Y-shaped dust cloud at $\sim 8''$ west of PGMW 3223. The line profiles are clearly split into 287 and

306 km s $^{-1}$ over an extent of $\sim 52''$ (or 13 pc) around PGMW 3223, with the Y-shaped dust cloud situated on the western edge. The echelle spectrum of the Y-shaped dust cloud also shows an embedded star with bright H α emission. This H α -emission star, marked as “H α star” in Figures 1c and 9, is probably a newly formed massive star.

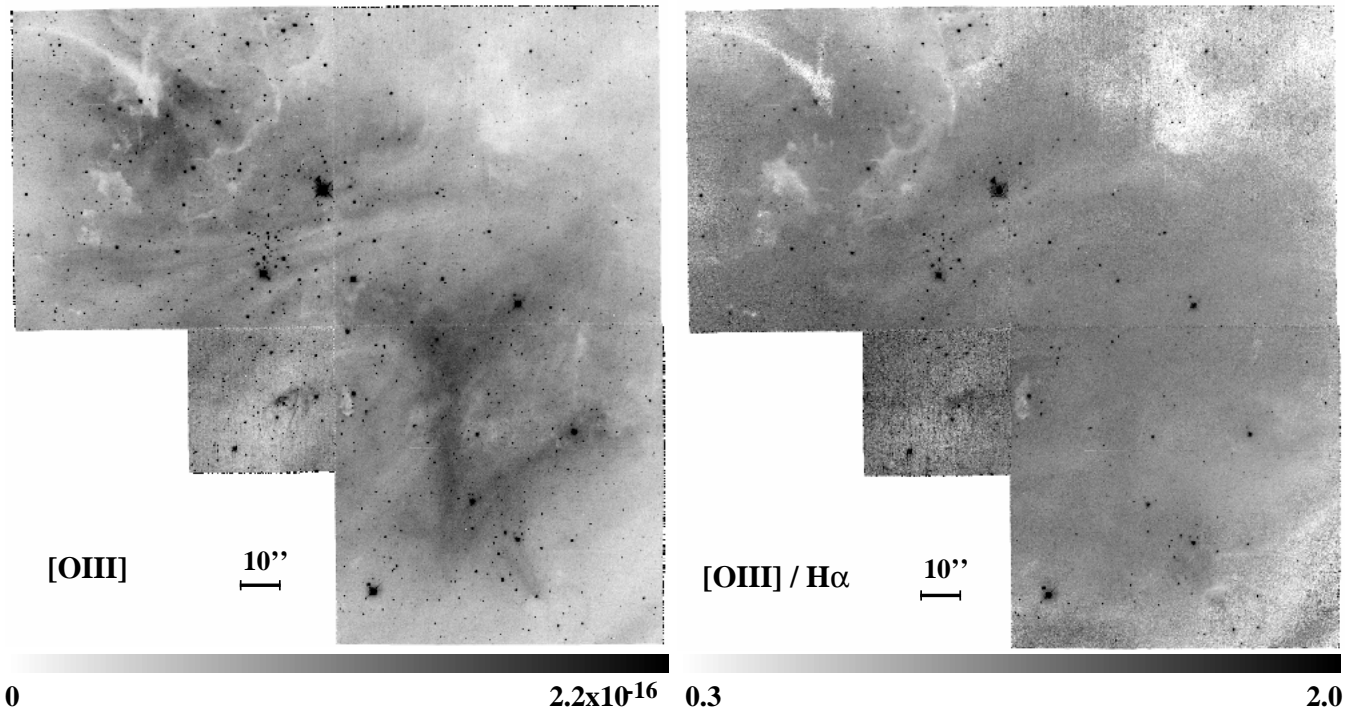


FIG. 5.—Same as Fig. 4, for N180B

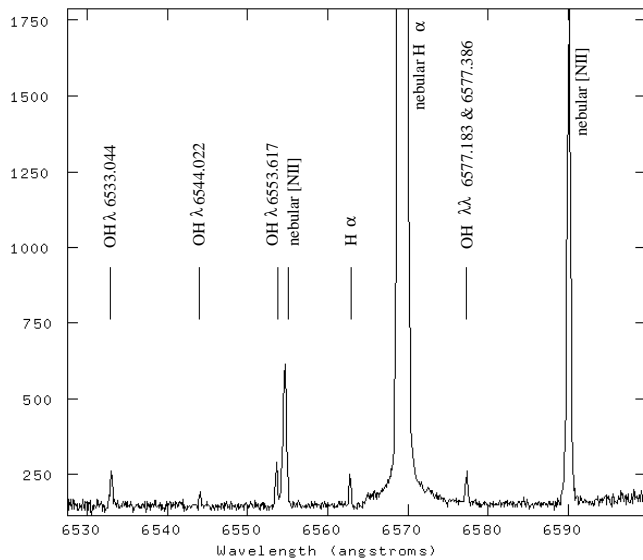


FIG. 6.—Identification of the telluric and nebular lines present in the echelle spectra of N11B and N180B. The telluric lines are narrow and unresolved. The intensity scale is arbitrary.

Similar velocities are also recorded at $67''$ – $86''$ west of PGMW 3223, corresponding to extensions of the bright H π region. Everywhere else, the line profiles are purely Gaussian and are centered at $\sim 296 \text{ km s}^{-1}$.

8. East-west and north-south of PGMW 3120 (O5.5 V(f^*)). These two slit positions are selected to study the ring nebula around PGMW 3120. As seen in the WFPC2 H α image (Figs. 1c and 2), this ring nebula is incomplete with a radius of $3''$ – $6''$. Both the H α and [N π] lines are clearly split, with the two components at 281 and 306 km s^{-1} . However, the spatial extent of the region with line splitting is from $12''$ east to $26''$ west of the star and from $7''$ north to at least $40''$ south of the star. This expanding shell is much more extended than the morphologically identified ring nebula, and it encompasses the nearby O7 V star PGMW 3102 (see Figs. 9 and 11). The small H α arc structure is probably fortuitous. At $40''$ – $54''$ south of PGMW 3120, high-velocity gas at $\sim 347 \text{ km s}^{-1}$ is seen. This region is within Rosado et al.'s (1996) region d, and the velocity features are similar to those within region d sampled by the north-south echelle slits centered on PGMW 3058 and PGMW 3053. The eastern part of the east-west slit samples

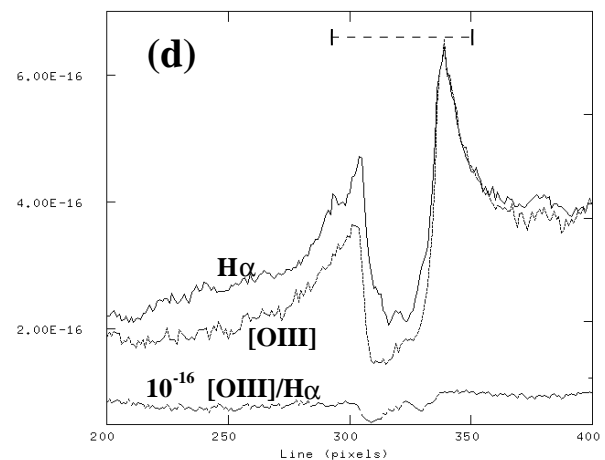
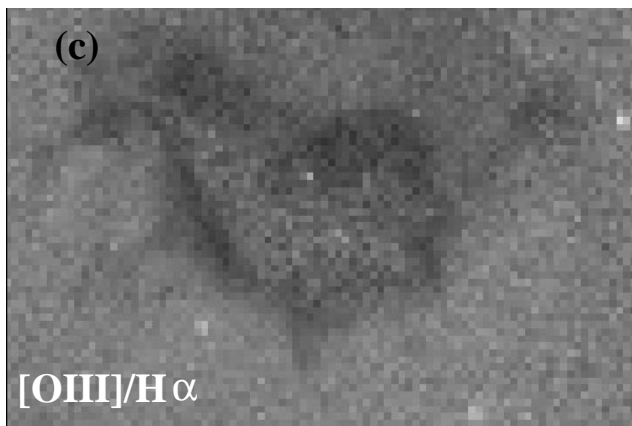
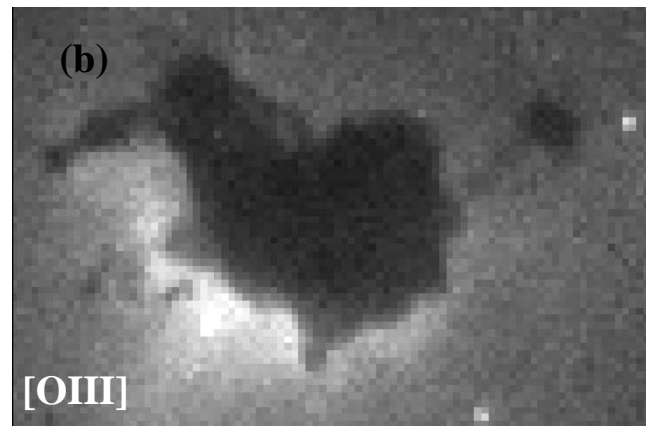
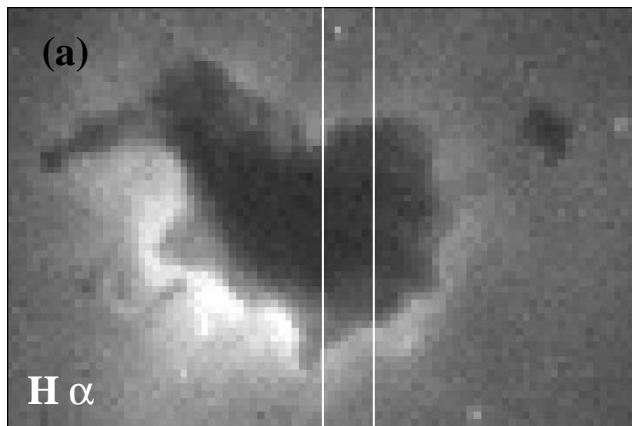


FIG. 7.—(a): HST WFPC2 H α . (b): [O III]. (c): [O III]/H α ratio map of the kiwi-shaped dust cloud in N11B. The size of each image is $9'' \times 6''$. Note that these images are presented with emission in white, unlike the other figures. (d): Profiles of H α and [O III] surface brightness, and the [O III]/H α ratio (multiplied by 10^{-16}) along a cut bounded by the white lines on the H α image in the top left panel. The X-axis is in units of WFC pixel with a size of $0''.1$. The profiles plot extends further than the images shown in (a), (b), and (c); the dashed line above the profiles marks the $6''$ extent.

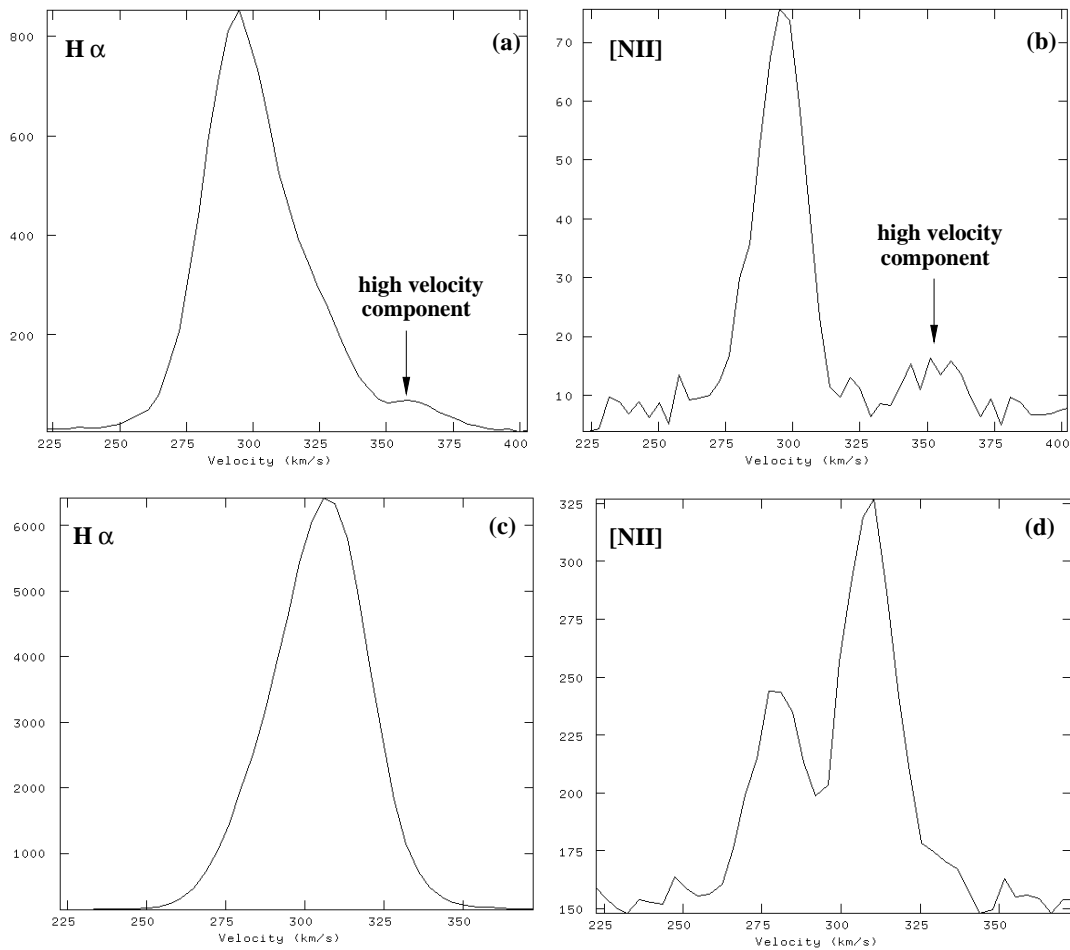


FIG. 8.—Examples of H α and [N II] line profiles. (a) H α and (b) [N II] profiles at 47°5 south of PGMW 3053. (c) H α and (d) [N II] profiles around PGMW 3120.

part of Rosado et al.'s (1996) region a. Near stars PGMW 3204 and 3209, the H α line is split into 287 and 308 km s⁻¹, similar to those seen in the observation along the east-west slit centered on PGMW 3204.

3.1.4. Integrated View of N11B

The analysis of the echelle observations of N11B shows clearly that expanding shells and high-velocity, accelerated material exist in this young H II region. Most of the expanding shells do not show identifiable morphological counterparts, while the morphologically identified ring nebulae do not show expanding shell structures.

Most of the expanding shells encompass groups of massive stars; only a few surround single massive stars. Some irregular expanding features with expansion velocities 10–15 km s⁻¹ are detected, but no clear association with massive stars can be identified. The expanding shells and high-velocity features associated with recognizable massive stars are described below and summarized in Table 3. The spectral types and expected wind luminosity, calculated using the wind velocities from Prinja et al. (1990) and the mass-loss rates from de Jager et al. (1988), are also given in this table.

The shell encompassing the large concentration of massive stars at the southwest corner of N11B corresponds to Rosado et al.'s region d. This shell has the largest expan-

sion velocity, ~ 20 km s⁻¹, and the largest dimension, ≥ 20 pc (see the velocity plots of north-south of PGMW 3053, 3058, and 3120 in Fig. 9). The south boundary of this shell is unknown, as it is not covered in our echelle slits. The second largest shell corresponds to Rosado et al.'s region a. This shell encompasses the concentration of stars including PGMW 3204, 3209, and 3223. Its expansion velocity is ~ 10 km s⁻¹ and its diameter is ~ 14 pc (see, e.g., the velocity plot of east-west of PGMW 3204 in Fig. 9).

One expanding shell might be associated with a single star, the uncataloged bright star between PGMW 3223 and PGMW 3224. This shell has an expansion velocity of ~ 12 km s⁻¹, and a diameter of ~ 10 pc (see the velocity plot of PGMW 3168 in Fig. 9). A high-velocity feature, redshifted by ~ 60 km s⁻¹ and extending over ~ 7 pc, is detected near PGMW 3061, an O3 III(f*) star (see the velocity plots of north-south of PGMW 3053 and PGMW 3058 in Fig. 9). It is possible that PGMW 3061 is responsible for its acceleration. The most interesting expanding structure around a single massive star belongs to PGMW 3160, an early K I supergiant. An expanding blister structure, with a size of ~ 5 pc and expansion velocity of ~ 15 km s⁻¹, is detected in the vicinity of this red supergiant (see the velocity plot of north-south of PGMW 3157). This expanding structure is most likely produced by the O star progenitor of PGMW 3160 during the main-sequence stage.

Three characteristics appear to be common among the expanding shells in N11B. First, shells encompassing small

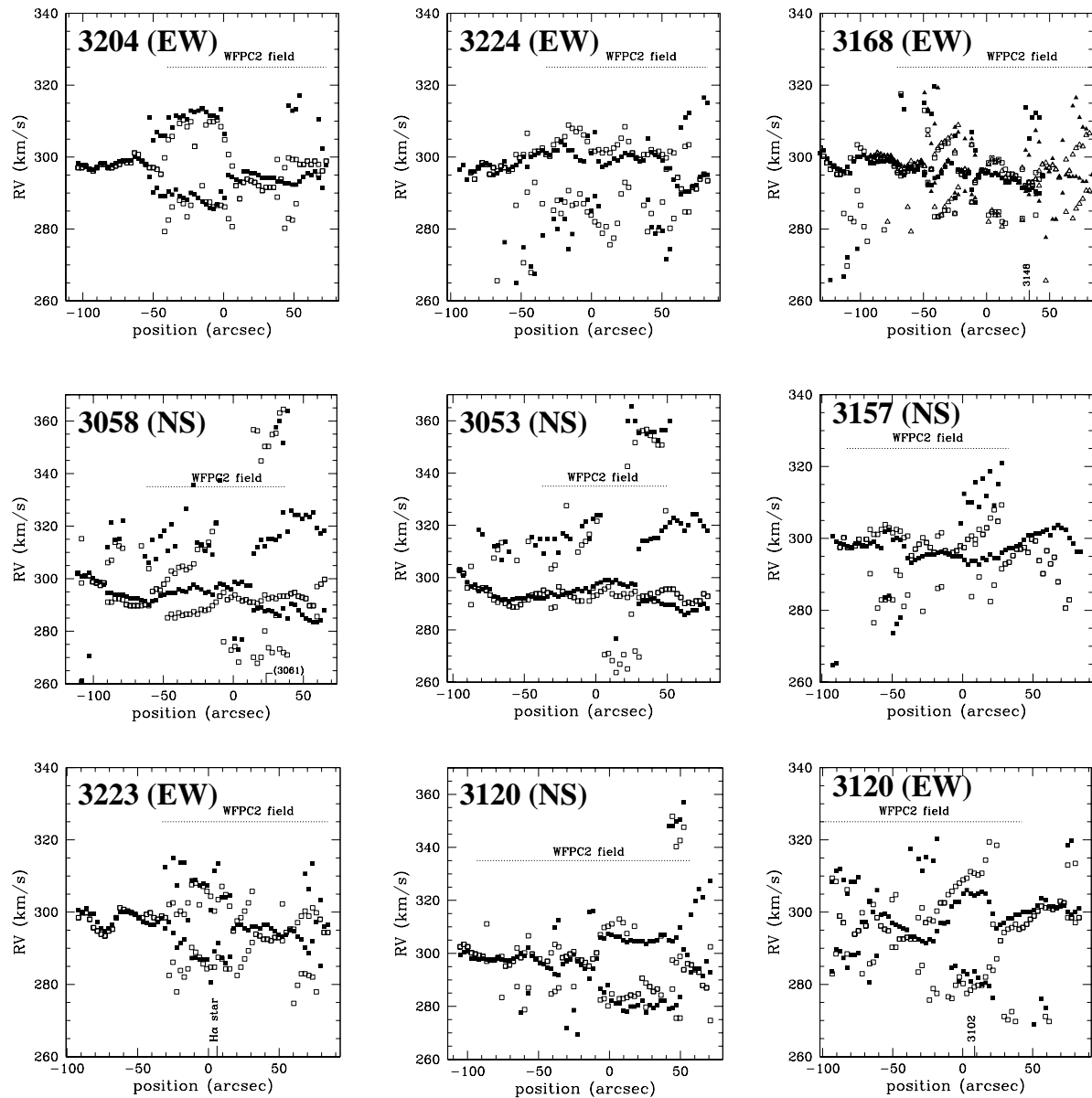


FIG. 9.—Radial velocity–position plots of the echelle slit positions in N11B for both $H\alpha$ and $[N\ II]$ lines. The radial velocities are heliocentric. The $H\alpha$ velocity components are plotted in filled symbols, and the $[N\ II]$ in open ones. The position axis is defined with the central star (labeled with its PGMW number at the upper left corner) at the origin ($0''$) and increasing toward south or west. When another star is present in the spectrum, its position is marked by a short vertical line labeled with its PGMW number. Stars close to the slit position are also marked, but in parentheses. The dashed line in each plot indicates the range covered by the WFPC2 images. The two observations east-west of PGMW 3168 were not centered at the same position; they were analyzed individually and plotted in different symbols (*squares and triangles*).

numbers of massive stars have sizes of 10–15 pc and expansion velocities of $10\text{--}15\text{ km s}^{-1}$, barely supersonic for an isothermal sound velocity of 10 km s^{-1} at 10^4 K . Only the shell encompassing a large number of stars is larger and expands faster. Second, most of the expanding shells have the massive stars displaced from the shell centers, indicating a blister-like structure. Finally, some shells might be merging together, causing the boundaries between them to become unclear.

3.2. N180B

3.2.1. Surface Brightness and Density

The most remarkable feature in the WFPC2 $H\alpha$ image of N180B is the prevalence of dust clouds in different sizes and shapes. There are 30 pc long dust streamers running across

the face of the $H\ II$ region, prominent antenna-shaped dust cloud at the northeastern corner of the field, large patches of dust clouds at the northwestern and southwestern corners of the field, and Bok-globule-like small dust clouds embedded in ionized gas near the large dust clouds. These dusty features amply portray a young star formation region.

The $H\alpha$ surface brightness is heavily modulated by the large-scale dusty features. The brightest emission region in N180B is located in the southwestern quadrant of N180B. Its V-shaped rim is dotted with four bright O and B stars: MGSD 152, 184, 187, and 197 (see Fig. 1d for the positions and spectral types). As shown below, this higher surface brightness is caused by a higher density in this region.

Again, we have adopted a simplistic two-density-component model for the structure of N180B: a global

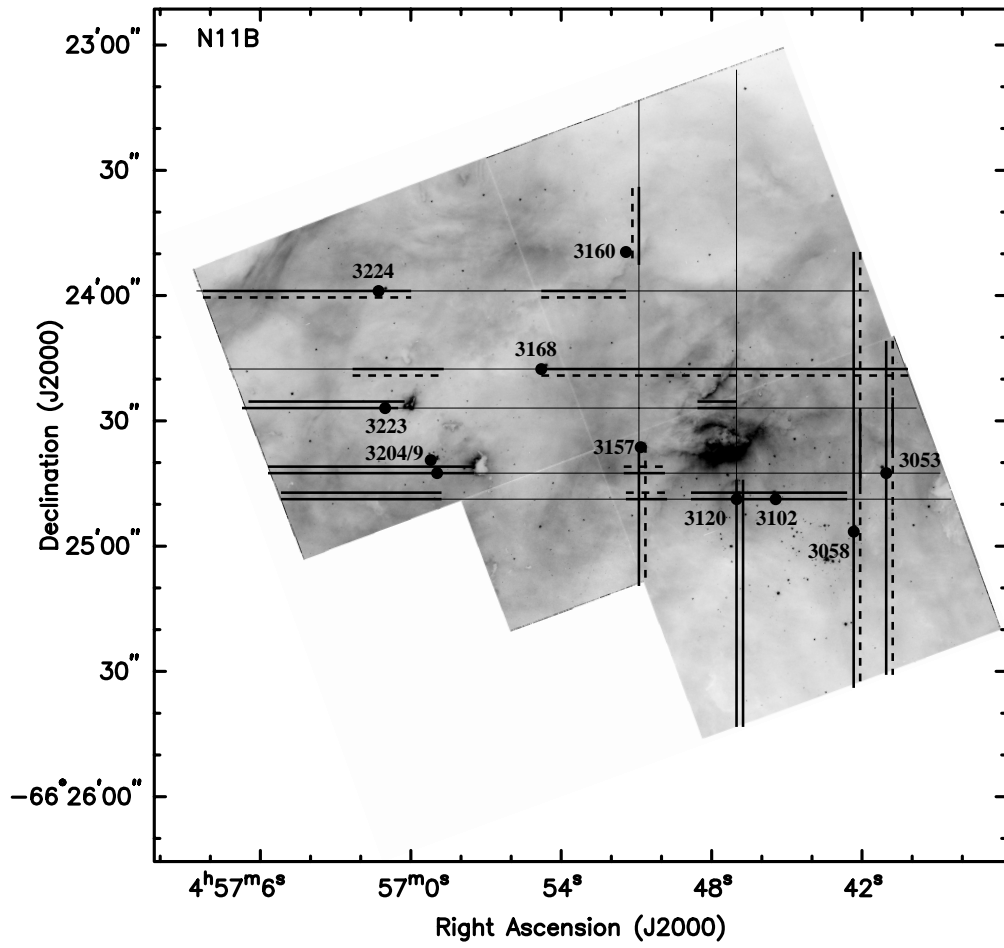


FIG. 10.—*HST* WFPC2 $H\alpha$ image of N11B superposed by spectral features: a thin line for a Gaussian profile, thick double lines for lines with clear splitting, and thick solid-dashed lines for asymmetric profiles indicating additional faint components. The relevant reference stars are also labeled. Note that the widths of the lines do not represent the slit width, which is $1''.65$.

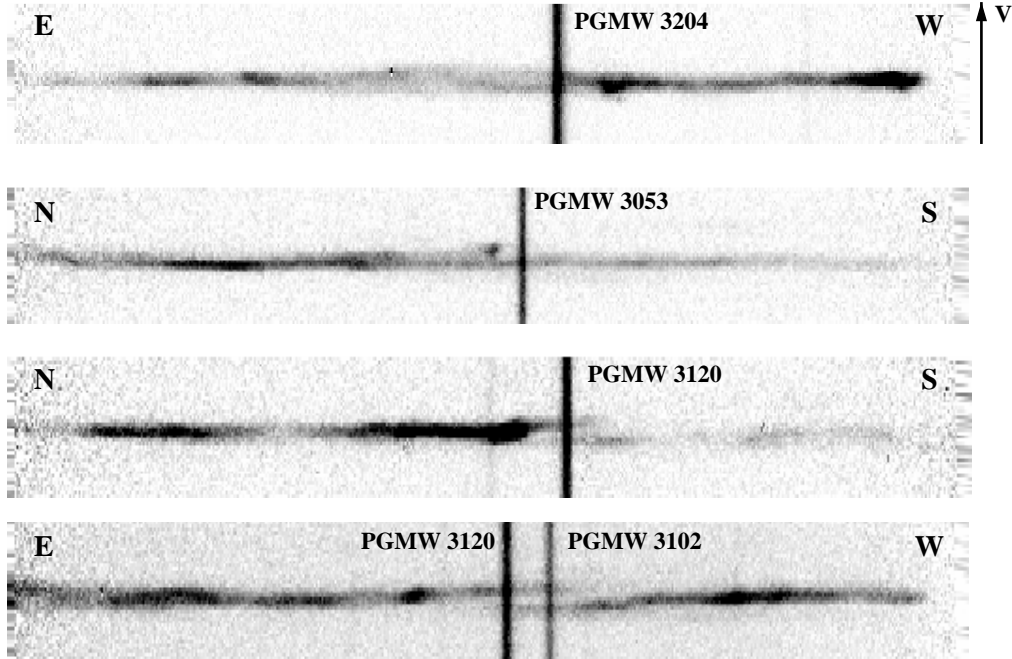


FIG. 11.—Four examples of echellograms of the [N II] line. The slit orientation and stars along the slit are marked. Each panel is $185''$ along the horizontal direction and 200 km s^{-1} along the vertical direction. The beginning of the splitting region seen to the north of PGMW 3053 is clearly marked by a bright knot in the spectrum. The line splitting around PGMW 3120 is clearly visible, while the “line-splitting” region to the east of PGMW 3204 shows only a broad line.

TABLE 3
SUMMARY OF EXPANDING SHELLS PROPERTIES

ID	Stellar Content	L_w (stars) (10^{36} ergs s $^{-1}$)	Diam. (pc)	V_{exp} (km s $^{-1}$)	t_6 (10^6 yr)	L_w (model) (10^{36} ergs s $^{-1}$)
N11B-1	All stars at the southwest of N11B	...	$\geq 20^a$	20
N11B-2	PGMW 3204 (O6–7 V)	1.03	14	10	0.41	0.25
	PGMW 3209A (O3 III(f*)) ^b	15.8
	PGMW 3209B (O9 V)	0.14
	PGMW 3209C (O7 V)	0.59
	PGMW 3209D (O9 V)	0.14
	PGMW 3209E (O9.5 V)	0.03
	PGMW 3209F (O9.5 V)	0.03
	PGMW 3223 (O8.5 V)	0.08
N11B-3	Uncataloged star between PGMW 3223 and 3224	...	10	12	0.25	0.22
N11B-4	PGMW 3160 (early KI)	...	5	15	0.10 ^c	0.11
N180B-1	MGSD 214 (O3–4)	13.4	22	20	0.33	2.97

^a The insufficient spatial coverage of our echelle spectra to the south of N11B does not allow us to assign accurate boundaries to this expanding shell.

^b The spectral types of components A–F are taken from Walborn et al. (1999)

^c As the star is evolved, the bubble dynamics must have been modified; therefore these numbers may not be meaningful.

uniform component and local density enhancements. We measured the surface brightness of a featureless region at 3".5 north and 17" west of the star MGSD 119 to represent the global component. Adopting a diameter of 225" (or 56 pc) for N180B and an $E(B-V)$ of 0.12 mag (MGSD), we derive a rms density of ~ 9.5 cm $^{-3}$. We have measured the surface brightness of the bright, V-shaped emission region and the arc around MGSD 218. The rms densities of these two regions are a few times 10 cm $^{-3}$. The results are detailed in Table 2.

3.2.2. Morphology and Excitation

Of all stars in the WFPC2 H α image of N180B, only two stars, MGSD 174 and MGSD 218, are surrounded by small (5"–10") arcs or circular filaments that can be identified as ring nebulae. No spectral classification was available for these two stars, but the UBV photometric data and extinction reported by MGSD suggest that the spectral type of MGSD 174 is B2 V and MGSD 218 B5 V. A few stars may be associated with faint nebular features, such as a small circular emission region around MGSD 82 and the "comet tail" of MGSD 106; however, these morphologies do not suggest a bubble structure.

The average [O III]/H α ratio of N180B is high, ~ 0.9 , comparable to that of N11B. In the [O III]/H α ratio map of N180B, most dust clouds, but not the long dusty streamers, are seen as regions with low [O III]/H α ratio, ~ 0.4 . The dusty streamers across the face of the H II region show [O III]/H α ratios similar to that of the overall H II region itself. The [O III]/H α ratio map of N180B appears much more uniform than that of N11B. Again, no sharp, shock-excited features are identified.

The small dust cloud near MGSD 168 has [O III]/H α ratios of 0.6–0.9, which are higher than those in N11B, such as the kiwi-shaped dust cloud. This difference is probably caused by the contamination of background/foreground emission. MGSD 168 is a mid-B star, so the H α emission from its neighboring dust cloud is weak compared with the background H α emission from gas ionized by early O stars, while the kiwi-shaped dust cloud in N11B is ionized by early O stars and its emission is much brighter than the background emission. Therefore the low [O III]/H α ratio in

the ionization front can be more easily seen in dust clouds near early O stars than later O or B stars.

The arc around MGSD 218 has a low [O III]/H α ratio, 0.45, indicating a low excitation at an ionization front. This kind of arc nebulae is frequently seen around early B main-sequence stars near dust clouds, e.g., LSS 3027 in the Galaxy (Chu 1983). MGSD 218 is a mid-B main-sequence star and could be luminous enough to produce a small H II region. The filamentary nebula around MGSD 174 may also have a low [O III]/H α ratio, but the signal-to-noise ratio of this region (imaged by the Planetary Camera of WFPC2) is too low to confirm the low [O III]/H α ratio.

Many nebular morphological features "disappear" into the background in the [O III]/H α ratio map. The small "comet tail" nebulosity of MGSD 106 and the large (20 pc in size), bright, V-shaped emission region both show [O III]/H α ratios similar to the average ratio of the entire H II region. It is interesting to note that MGSD 118 (O4 III(f*) star) is surrounded by a small emission region (size = 3") with a very high [O III]/H α ratio (~ 1.5), but this small nebula does not show any peculiar kinematic properties (see § 3.2.3).

3.2.3. Kinematic Properties

The kinematic structure of N180B is much more quiescent than that of N11B. Line splitting is seen only in the [N II] line within one expanding shell centered on MGSD 214 (O3–4). The H α line is narrow with an observed FWHM of ~ 30 km s $^{-1}$. Quadratically subtracting the instrumental FWHM of 13 km s $^{-1}$ and a thermal FWHM of 21 km s $^{-1}$ (assuming 10^4 K) from the observed width, we obtain a turbulent FWHM of only 17 km s $^{-1}$.

Figure 12 shows the velocity components of the H α and [N II] lines along each slit. Regions with more than one velocity component are marked on the H α image in Figure 13. The kinematic features along each slit position are described below.

1. North-south of MGSD 118 (O4 III(f*)). The northern part of this slit, including the location of MGSD 118, show narrow H α and [N II] lines centered at 245–250 km s $^{-1}$. From 28" south of MGSD 118 southward, the H α line shows an additional faint blue component, which has an

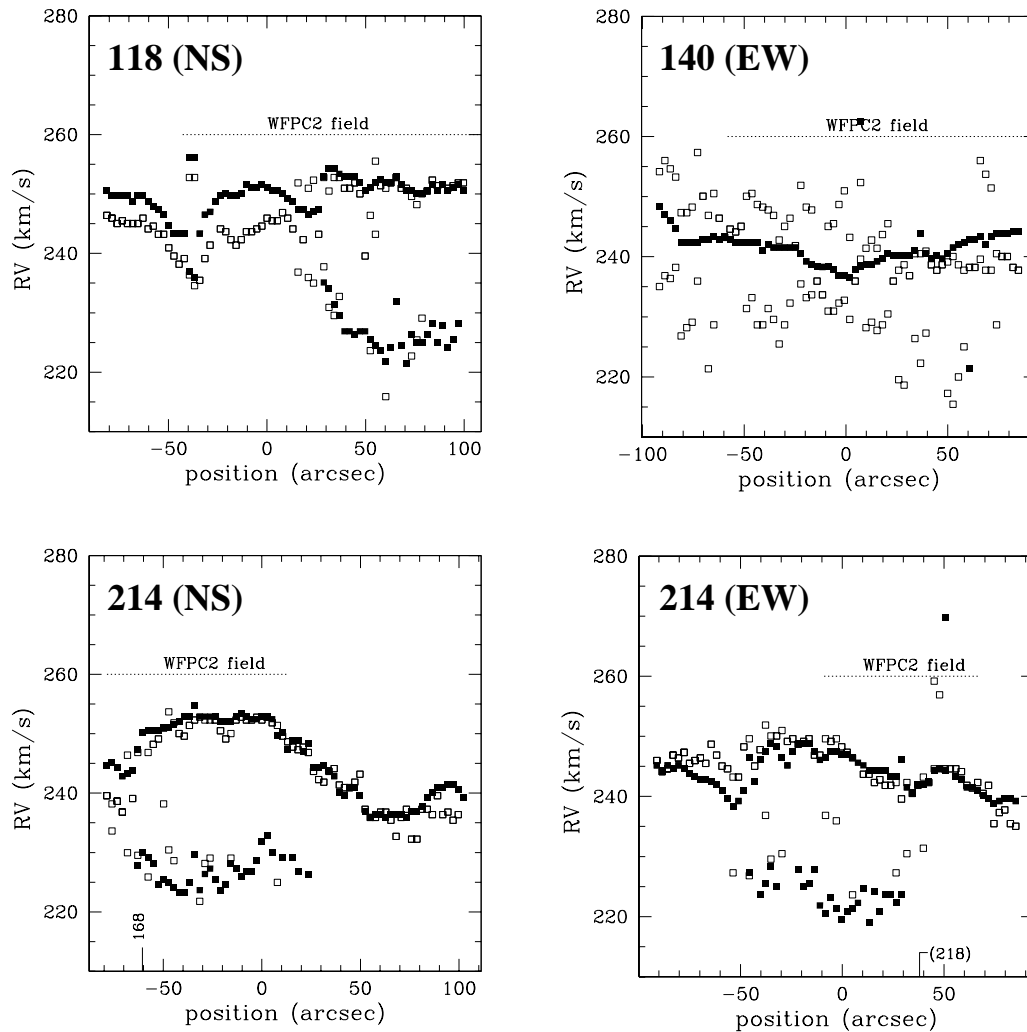


FIG. 12.—Same as Fig. 9, but for N180B. The star numbers are from MGSD.

extreme velocity at $220\text{--}225\text{ km s}^{-1}$. This suggests an expanding “blister,” with the far side roughly at the H II region’s systemic velocity and the near side expanding toward us with an expansion velocity of $\sim 20\text{ km s}^{-1}$. The northern edge of the region appears to be bordered by a broad arclike structure.

2. East-west of MGSD 140 (O3–4(f*)). Along the entire slit, the H α line appears to be a perfect Gaussian profile centered at $\sim 240\text{ km s}^{-1}$, while the [N II] line shows asymmetric line profiles indicating the existence of at least two velocity components. Spectral fits of the [N II] line yield components separated by $\leq 20\text{ km s}^{-1}$. The [N II] line is weak and noisy, and thus the velocity components show a large velocity scatter.

3. East-west and north-south of MGSD 214 (O3–4). From $50''$ east to $40''$ west of MGSD 214, both H α and [N II] lines show a bright component at $240\text{--}245\text{ km s}^{-1}$, near the H II region’s systemic velocity and a faint blue component at $220\text{--}225\text{ km s}^{-1}$. This velocity pattern suggests a blister expanding toward us. The north-south slit position shows that this expanding blister extends from $\sim 65''$ north to $\sim 25''$ south of MGSD 214. Outside this expanding blister region, the H α and [N II] lines are both narrow. Velocity gradients are seen along both slits, from one side of the expanding blister to the other side. It is not

clear whether the velocity variation of the bright component is attributed largely to the velocity gradient in the H II region or an expansion velocity variation in the receding side of the blister.

3.2.4. Integrated View of N180B

N180B is much more quiescent than N11B. Only one expanding blister is unambiguously detected around the O3–4 star MGSD 214. This blister is $\sim 22\text{ pc}$ in diameter and expanding away from the dense cloud at $\sim 20\text{ km s}^{-1}$ (see Table 3). No morphological counterpart of this expanding blister can be confidently identified. Two other stars have similar spectral types, MGSD 118 of type O4 III f* and MGSD 140 of type O3–4 (f*), but do not show unambiguous expanding shells or blisters around them. The FWHM of the H α line in the vicinity of these two stars suggests a turbulent component of $\leq 17\text{ km s}^{-1}$ FWHM, while the [N II] line shows additional faint blue components. It is possible that these two massive stars have blisters expanding at $\leq 10\text{ km s}^{-1}$.

The overall kinematic structure of N180B exhibits two characteristics. First, the H II region itself has a velocity gradient or variation at a level of $\sim 10\text{ km s}^{-1}$ across the nebula. Second, the expanding shells have a blister structure

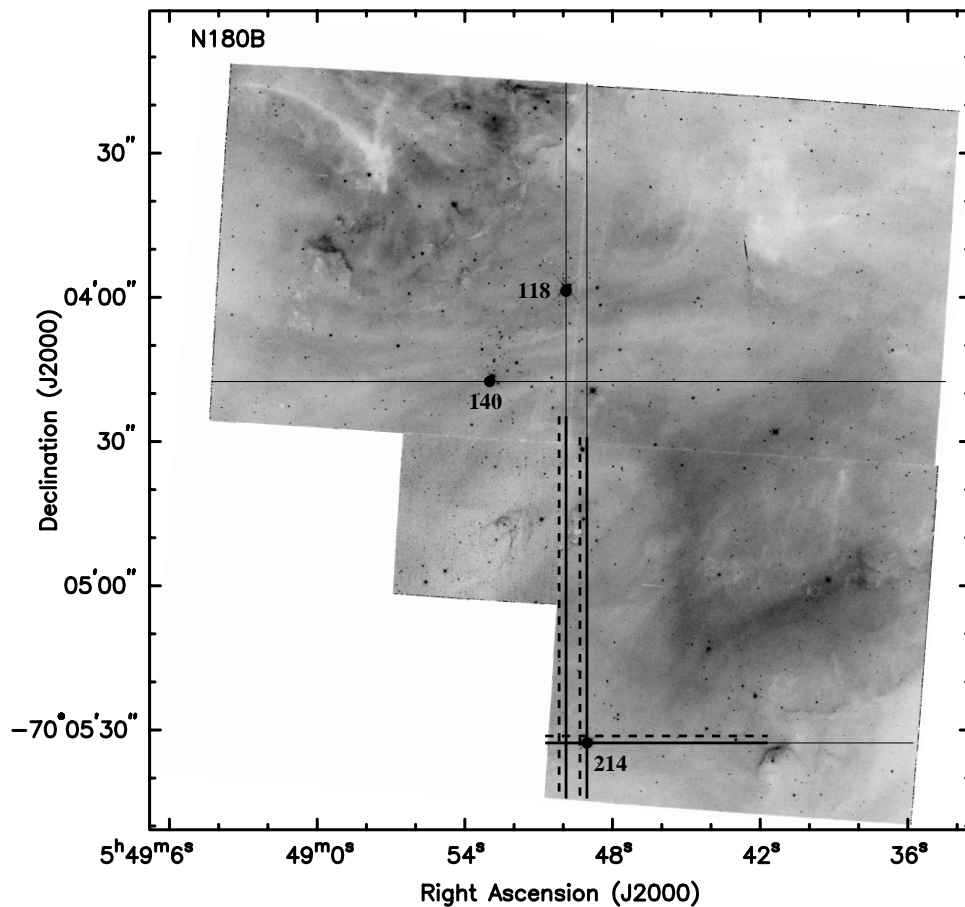


FIG. 13.—Same as Fig. 10, but for N180B. The star numbers are from MGSD.

expanding toward us, with the back side being the dense H II region itself.

4. DISCUSSION AND CONCLUSIONS

Massive stars are expected to produce wind-blown bubbles in the interstellar medium (Castor, Weaver, & McCray 1975; Weaver et al. 1977); however, main-sequence O stars are rarely seen surrounded by ring nebulae (Chu 1991). To search for wind-blown bubbles around main-sequence O stars, we have obtained high-resolution *HST* WFPC2 images and high-dispersion echelle spectra of two pristine H II regions, N11B and N180B, ionized by OB associations. The presence of O3 stars in these OB associations suggests that the H II regions are young and have not hosted any supernova explosion.

Analyzing the nebular morphologies and kinematics, we find that main-sequence O stars do blow bubbles in H II regions and that the bubbles can be detected kinematically¹¹ but not morphologically. Most of these main-sequence bubbles have expansion velocities of 10–15 km s⁻¹, which is comparable to or slightly larger than the isothermal sound velocity of ionized gas at 10⁴ K. The expanding bubbles generate only weak interstellar shocks as they expand into the ambient interstellar medium; therefore, no strong compression of the interstellar medium is

expected in these bubbles. Lacking sharp filaments in H α images, the slowly expanding bubbles blown by main-sequence stars cannot be identified morphologically. Ironically, in N11B and N180B the morphological features that are conventionally identified as ring nebulae, such as a bright nebular arc around a star, are not expanding shells (e.g., MGSD 218 in N180B or PGMW 3120 in N11B).

Three types of wind-blown bubbles are seen in those young H II regions: (1) bubbles around concentrations of massive stars within which the projected separation between neighboring stars is typically ≤ 1 pc, (2) bubbles around individual O stars if the distance to neighboring stars is large enough to maintain the individuality of the bubble, and (3) a fossil bubble around an evolved red supergiant that is isolated. It is conceivable that the formation of wind-blown bubbles or superbubbles in an H II region depends crucially on the spatial distribution of the stars. Tight concentrations of massive stars are conducive to the formation of superbubbles even before any supernova explosion. Wide separations between massive stars are conducive to the formation of single-star, main-sequence bubbles, and a bubble may remain distinct even after the star has evolved into a red supergiant.

The observed expansion dynamics of the interstellar bubbles can be compared with those expected in bubble models of Castor et al. (1975) and Weaver et al. (1977). With the observed radius (R in units of parsecs) and expansion velocity (V in units of kilometers per second), we can compute the dynamical timescale (t_6 in units of 10⁶ years) using $t_6 = 0.6R/V$ and further determine the ratio of stellar

¹¹ Note that the photoevaporation of dense knots will produce line broadening (see simulations by Brandner et al. 2000), not the expanding shells we observed.

wind luminosity (L_{36} in units of 10^{36} ergs per second) to ambient density (n_0 in units of H atoms per cubic centimeter) using $L_{36}/n_0 = (R/27)^5 t_6^{-3}$. We adopt the rms densities we have determined for N11B and N180B, 15 H atoms cm^{-3} and 9.5 H atoms cm^{-3} , respectively. We have listed in Table 3 the observed diameter and expansion velocity of the expanding shells, together with the derived dynamical timescale and wind luminosity. We have also listed the early-type stars encompassed within each shell and the stellar wind luminosity expected from these stars.

We will focus on the two best-defined shells in our sample. First we consider the bubble blown by the O3–4 star MGSD 214 in N180B. With a radius of 11 pc and an expansion velocity of $\sim 20 \text{ km s}^{-1}$, the bubble around MGSD 214 has a dynamical timescale of 3.3×10^5 yr. Adopting N180B's rms density of $\sim 9.5 \text{ cm}^{-3}$, we find a wind luminosity of $\sim 3 \times 10^{36} \text{ ergs s}^{-1}$ for MGSD 214. This is a factor of 10 lower than that expected from an O3–4 star. Next we consider the bubble blown by PGMW 3204, 3209, and 3223 in N11B. The star PGMW 3209, while dominated by an O3 III(†) star, is in fact composed of at least five additional late O stars (Walborn et al. 1999). For a radius of 7 pc and an expansion velocity of 10 km s^{-1} , the dynamical timescale of this bubble is 4.1×10^5 yr. Using a rms density of 15 cm^{-3} for N11B, we expect a wind luminosity of $\sim 2.5 \times 10^{35} \text{ ergs s}^{-1}$ for PGMW 3209, 2 orders of magnitude lower than that expected from an O3 III star. This discrepancy is even larger if we consider the additional stars in PGMW 3209 and the neighboring stars PGMW 3204 and PGMW 3223.

The large discrepancy between stellar wind luminosities derived from observations and models of bubbles and those expected from the actual stellar content is very interesting. A similar discrepancy has been reported for circumstellar bubbles (e.g., NGC 6888; García-Segura, & Mac Low 1995) and superbubbles (Oey 1996). Our results represent the first report for small interstellar bubbles. This discrepancy implies that either the observed stellar wind luminosities have been overestimated or Weaver et al.'s model cannot describe the observed bubbles. If the observed stellar wind luminosity is correct, our observed bubbles will be too small and expand too slowly. To reduce the discrepancy, we need to increase the ambient density, but it is unlikely that we have underestimated the rms density of the H II regions by factors of 10–100.

One further discrepancy underscores the deficiency of bubble models. The dynamical timescales of the bubbles in N11B, 0.1–0.5 Myr, are much shorter than the age of the

OB association LH10, 2–3 Myr (Walborn & Parker 1992). A similar discrepancy is also observed in superbubbles (Oey 1996). These problems need to be seriously considered in the future.

The content and distribution of wind-blown bubbles in N11B are different from those in N180B: N11B has a larger number of wind-blown bubbles and faster expanding bubbles. This difference is caused by the richness of massive stars in the encompassed OB associations. The OB association LH10 in N11B has 20 stars with stellar masses $M \geq 15 M_\odot$ (PGMW), while the OB association LH118 in N180B only has six (MGSD). Furthermore, LH10 has tighter concentrations of massive stars than LH118. Bubbles blown by concentrations of massive stars expand faster because of more powerful stellar winds. Therefore, N11B has more bubbles and the fastest expanding bubble is around its highest concentration of massive stars.

One property in common among the bubbles in N11B and N180B is the asymmetric structure—most of the bubbles are “blisters.” A face-on blister can be diagnosed from the asymmetry in line profiles, which show a bright component near the H II region's systemic velocity and a faint component blueshifted or redshifted to higher velocities. A blister viewed from the side can be diagnosed by the displacement of the responsible star from the projected bubble center. As the stars in N11B and N180B are still young and must be located close to their birth places, their blister-like bubbles suggest that the stars are located near the surface of a dense cloud, where a density gradient is present; as opposed to embedded in the deep interior of a cloud, where the ambient density might be more uniform.

Finally, we note that numerous small dust clouds resembling Bok globules and elephant trunks are detected in N11B and N180B. If the ambient interstellar gas pressure is high enough, star formation might be triggered in these small dust clouds. At least one such cloud in N11B, the Y-shaped dust cloud near PGMW 3223, shows signs of on-going star formation at its tip. An infrared survey of N11B and N180B would reveal whether star formation has started in other small dust clouds.

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